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A Positive Energy Theorem for Asymptotically deSitter Spacetimes

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ABSTRACT

We construct a set of conserved charges for asymptotically deSitter spacetimes that correspond to asymptotic *conformal isometries*. The charges are given by boundary integrals at spatial infinity in the flat cosmological slicing of deSitter. Using a spinor construction, we show that the charge associated with conformal time translations is necessarily positive and hence may provide a useful definition of energy for these spacetimes. A similar spinor construction shows that the charge associated with the time translation Killing vector of deSitter in static coordinates has both positive and negative definite contributions. For Schwarzschild-deSitter the conformal energy we define is given by the mass parameter times the cosmological scale factor. The time dependence of the charge is a consequence of a non-zero flux of the corresponding conserved current at spatial infinity. For small perturbations of deSitter, the charge is given by the total comoving mass density.

1 Introduction

In this work we will study the notion of mass in asymptotically deSitter spacetimes. Recent cosmological observations indicate that our universe is best described by a spatially flat, Friedmann-Robertson-Walker cosmology with the largest contribution to the energy density coming from some form of ‘dark energy’¹, matter with an approximate equation of state $p = w\rho$ with $-1 \leq w \leq -1/2$. One possibility for the dark energy is a positive cosmological constant, which is $w = -1$. In this case, as the universe expands the cosmological constant will become increasingly dominant, so that in the future our spacetime will tend to increasingly approximate a region of deSitter spacetime. These observations have led to considerable renewed interest in all things deSitter, including both the classical and quantum mechanical properties of asymptotically deSitter spacetimes.

Our particular focus in this paper will be to demonstrate positivity properties for the mass of asymptotically deSitter spacetimes. At first glance, there are at least three lines of reasoning that suggest that such positivity properties should not exist for deSitter. First, Witten’s proof of the positive energy theorem for asymptotically flat spacetimes [2] and similar results for asymptotically anti-deSitter spacetimes [3][4][5] rely on flat Minkowski spacetime and anti-deSitter spacetime respectively being supersymmetric vacuum states. DeSitter spacetime is famously not a supersymmetric vacuum state of gauged supergravity. Therefore, one would not expect to be able to prove a positive energy theorem. Second, in global coordinates the spatial sections of deSitter are closed 3-spheres. These slices have no spatial infinity at which to define either asymptotic conditions on the metric, or an ADM-like expression for the mass. Rather, as mentioned above for our universe, asymptotic conditions are defined in the past and/or the future. Finally, Abbott and Deser have suggested a definition of mass for asymptotically deSitter spacetimes [6]. However, they argued that this mass should receive negative contributions from fluctuations outside the deSitter horizon.

Nonetheless, one further line of reasoning suggests that a positive energy theorem should exist for deSitter, and we will see that this is in fact the case. The positive energy theorem for asymptotically flat spacetimes may be generalized to Einstein-Maxwell theory [7] with the result that, if M and Q are the ADM mass and total charge of the spacetime, then $M \geq |Q|$. This bound is saturated by the $M = |Q|$ Reissner-Nordstrom spacetimes and more generally by the well known MP spacetimes [8][9], which describe collections of $M = |Q|$ black holes in mechanical equilibrium with one another. This bound has a natural setting in $N = 2$ supergravity, and in this context the MP solutions can be shown to preserve 1/2 the supersymmetry of the background flat spacetime [7]. These results closely parallel those associated with the BPS bound for magnetic monopoles in Yang-Mills-Higgs theory [10]. Hence, the positive energy bounds in general relativity are regarded as gravitational analogues of BPS bounds.

It is less well known that multi-black hole solutions also exist in a deSitter back-

¹See *e.g.* [1] for a recent summary

ground [11]. The single object solution in this case is a $M = |Q|$ Reissner-Nordstrom-deSitter (RNdS) black hole. These spacetimes reduce to the MP spacetimes if one sets the cosmological constant $\Lambda = 0$, and it seems reasonable to call them MPdS spacetimes². It is plausible that the MPdS spacetimes saturate a BPS bound in the form of a positive energy theorem for asymptotically deSitter spacetimes.

It is interesting to ask how the positive energy theorem we present overcomes the three objections stated above? First, with regard to supersymmetry, we will see that our construction is related to the deSitter supergravity theory presented in [12]. The super-covariant derivative operator acting on spinors, which we define below, coincides with the differential operator in the supersymmetry transformation law for the gravitino field in [12]. Although the quantized deSitter supergravity theory is non-unitary, we see that it is nonetheless useful for deriving classical results.

Second, we must specify an asymptotic region where the mass is to be defined. As stated above, an asymptotically deSitter spacetime is specified by conditions at past and/or future infinity. To define a mass for such spacetimes, we must consider spatial slices that asymptotically approach one of these regions. For exact deSitter spacetime, flat cosmological coordinates, as in equations (3) and (4) below, provide an example of such a slicing. Figure (1) below shows the conformal diagram for deSitter spacetime, with spatial slices at two different times sketched in. We see that the point at the upper left hand corner of the diagram can be reached by going to infinite distance along a spatial slice and can therefore be regarded as spatial infinity. We will assume below that the spacetimes under consideration have an asymptotically deSitter region in the future and consider a slicing such that the metric approaches that of deSitter spacetime in the flat cosmological slicing. We will adopt a pragmatic approach to fall-off conditions, requiring that our constructions be well defined, for example, for a galaxy in deSitter.

Finally, there is the issue of the non-positivity of the mass for asymptotically deSitter spacetimes defined by Abbott and Deser [6]. The construction in reference [6] holds generally for any class of spacetimes asymptotic, with appropriate fall-off conditions, to a fixed background spacetime. If the background spacetime has isometries, it is shown that a conserved charge, defined by a boundary integral in the asymptotic region, can be associated with each of the background Killing vectors. For example, the ADM mass of an asymptotically flat spacetime corresponds in this way to the time translation isometry of Minkowski spacetime. DeSitter spacetimes are maximally symmetric and therefore also have a maximal number of conserved charges. A natural choice to call the mass of an asymptotically deSitter spacetime is the charge associated with the time translation Killing vector for the deSitter background written in static coordinates [6]. However, this Killing vector becomes

²For $\Lambda > 0$ the physical properties of the MPdS spacetimes depart from those of the MP spacetimes in interesting ways. The relation $M = |Q|$ is no longer the extremality condition for RNdS black holes with $\Lambda > 0$. Rather, it can be shown to be the condition for thermal equilibrium between the black holes and the deSitter background. For $M = |Q|$ the Hawking temperatures associated with the black hole and deSitter horizons are equal. The solutions with $\Lambda > 0$ are also not static. The black holes follow the expansion or contraction of the background deSitter universe.

spacelike outside the deSitter horizon and correspondingly the mass receives negative contributions from matter or gravitational fluctuations outside the deSitter horizon [6][13]. Unlike anti-deSitter spacetimes, deSitter spacetimes have no globally timelike Killing vector fields and therefore one would not expect to find charges with positivity properties.

DeSitter spacetime does have a globally timelike conformal Killing vector fields (CKV). If (anti-)deSitter spacetime is viewed as a hyperboloid embedded in a flat spacetime of the correct signature, the CKVs are simply the projections onto the hyperboloid of the translation Killing vectors of the flat spacetime. For deSitter, time translation symmetry in the embedding spaces yields a globally timelike CKV. We will show that there is a conserved charge associated with this CKV, and that the charge is positive.

This positive, conserved charge, which we denote Q_ψ , has a simple interpretation in the perturbative limit. Let $\delta\rho$ be the perturbation to the mass density as measured by a comoving cosmological observer. We show that $\delta Q_\psi = a(t) \int \delta\rho$, *i.e.* for linear perturbations, the conformal mass at time t is the scale factor times the total comoving mass perturbation. The time dependence of the charge corresponds to a non-zero flux of the corresponding conserved current at spatial infinity. For Schwarzschild-deSitter one finds $Q_\psi = a(t)M$, where M is the constant mass parameter in the metric.

The subsequent sections of the paper are structured as follows. In section (2), we introduce some necessary elements of (anti-)deSitter geometry. In section (3) we present a positivity proof, in the manner of Witten [2], for the charge Q_ψ defined in terms of a boundary integral over spinor fields in the asymptotic deSitter region. In order to interpret this result, in section (4) we show that the Abbott and Deser construction [6] can be straightforwardly generalized to associate a conserved charge Q_ξ with *any* vector field. We speculate that this result is related to Wald's construction [14] of a Noether charge associated with an arbitrary spacetime diffeomorphism. We show that the charge Q_ψ constructed in section (3) is identical to the charge Q_ξ , where the vector ξ^a is a conformal Killing vector of the background deSitter spacetime. We will also discuss the dynamical role of this conformal energy as the Hamiltonian generating conformal time evolution in the sense of reference [15]. In an appendix we give a spinor construction that applies to the Killing vectors of deSitter and shows that the exact volume integrand for the charge has both positive and negative semi-definite contributions³.

In closing the introduction, we note a number of questions suggested by our results, which we will leave for future work. First is the extension to Einstein-Maxwell theory. We would like to check that the MPdS spacetimes of [11] do indeed saturate a BPS bound related to the positive energy theorem presented here. A second interesting direction would be to explore the relationship between the conserved charges Q_ξ of section (4) and Wald's Noether charges [14]. A third direction would be a more formally complete treatment of asymptotic falloff conditions and a calculation of the

³See, however, reference [16] where it is argued that the Abbott & Deser mass nonetheless turns out to be overall positive.

Poisson bracket algebra of conserved charges as in [15]. A fourth would be to see if our results are useful in the context of the dS/CFT conjecture (see [17] and references thereto).

Finally, we would like to comment on the relation of two papers to the present work. First is an earlier paper of our own [18]. In that work, we derived the basic positivity result for asymptotically deSitter spacetimes presented in section (3) of the present paper. However, we incorrectly interpreted the spinor boundary term as the charge associated with the time translation Killing vector of deSitter spacetime in static coordinates, *i.e.* the Abbott & Deser mass. The more complete presentation here demonstrates the correct interpretation in terms of a new conserved charge Q_ξ associated with translations in the conformal time coordinate. Second is a recent paper by Shiromizu et. al. [16]. This paper independently arrives at the association of the charge Q_ψ with the conformal time translation CKV⁴.

2 Some DeSitter Basics

In this section, we present a number of properties of deSitter spacetimes that will be useful in following sections. In particular, we will focus on the conformal Killing vectors of deSitter and expressions for them in terms of suitably defined Killing spinors. To keep the formulas simple, we will work in $D = 4$. Extension of the results to higher dimensions is straightforward. For comparison, we will also present many parallel formulas for anti-deSitter spacetimes.

Carving Out (A)dS from Flat Space

Start with the embeddings of 4-dimensional dS and AdS as hyperboloids respectively in $(4+1)$ and $(3+2)$ dimensional Minkowski spacetime. These can be studied together by writing the embedding equation for (A)dS with cosmological constant $\Lambda = 3\kappa/R^2$ as

$$-(X^0)^2 + (X^1)^2 + (X^2)^2 + (X^3)^2 + \kappa(X^4)^2 = R^2, \quad (1)$$

where $\kappa = +1$ for deSitter and $\kappa = -1$ for anti-deSitter. It will be useful below to write the five dimensional flat metric as

$$ds_5^2 = \kappa dR^2 + R^2 k_{ab} dx^a dx^b. \quad (2)$$

where x^a with $a, b = 0, 1, 2, 3$ are arbitrary coordinates on $(A)dS$ and $g_{ab} = R^2 k_{ab}(x^c)$ is the $(A)dS$ metric with radius R . We will also frequently choose flat cosmological coordinates (t, \vec{x}) for deSitter, in which the dS metric takes the form

$$ds^2 = -dt^2 + a(t)^2 d\vec{x} \cdot d\vec{x}, \quad a(t) = e^{Ht}, \quad (3)$$

⁴The overall line of reasoning and focus of reference [16] is quite different from the present paper. In particular positivity of the charge Q_ψ is derived starting from the ordinary positive energy theorem in a conformally related asymptotically flat spacetime.

where the Hubble constant $H = R^{-1}$. These cosmological coordinates are related to the flat embedding coordinates in equation (1) via the relations

$$t = R \log \left(\frac{X^0 + X^4}{R} \right), \quad x^k = \frac{RX^k}{X^0 + X^4}, \quad k = 1, 2, 3. \quad (4)$$

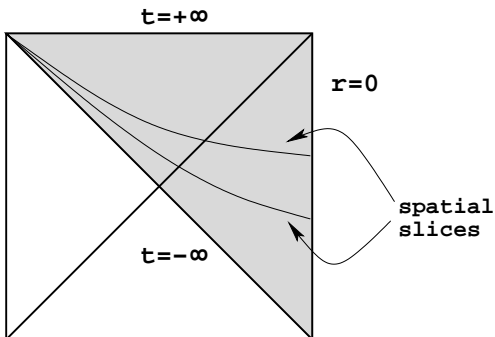


Figure 1: The conformal diagram for deSitter spacetime with the region covered by the flat cosmological coordinates of equations (3) and (4) shaded and spatial slices at two different times are indicated.

Conformal Killing Vectors

The isometry groups of dS and AdS are the Lorentz groups $SO(4, 1)$ and $SO(3, 2)$ of the respective flat embedding spacetimes. However, the groups of conformal isometries for both dS and AdS (as well as for $(3 + 1)$ dimensional Minkowski spacetime) are the same⁵ $SO(4, 2)$. One can think of the extra conformal isometries as arising from the translational Killing vectors of the flat 5 dimensional embedding spacetime. The CKV's of $(A)dS$ are simply the projections of these vectors onto the $(A)dS$ hyperboloids. This can be demonstrated easily starting from the form of the 5 dimensional flat metric in equation (2). Let $\xi^{(A)} = \partial/\partial X^A$. The vectors $\xi^{(A)}$ satisfy $\nabla_B^{(5)} \xi_C^{(A)} = 0$. After transforming to the (R, y^a) coordinates adapted to the hyperboloid, we then have in particular

$$0 = \nabla_a^{(5)} \xi_b^{(A)} \quad (5)$$

$$= \nabla_a \xi_b^{(A)} - \Gamma_{ab}^R \xi_R^A, \quad (6)$$

where ∇_α is the $(A)dS$ covariant derivative operator. The relevant Christoffel symbol $\Gamma_{\alpha\beta}^R$ is simply related to the extrinsic curvature of the embedding and is found from

⁵In a general the spaces $(A)dS_p$ share the conformal group $SO(p, 2)$ with $(p - 1, 1)$ dimensional Minkowski spacetime.

the metric (2) to be $\Gamma_{\alpha\beta}^R = -\frac{\kappa}{R}g_{\alpha\beta}$. For the $(A)dS$ space of radius R , we then have

$$\nabla_a \xi_b^{(A)} = -\frac{\kappa}{R}g_{ab}\xi_R^{(A)} \quad (7)$$

Projection of the five dimensional vectors $\xi^{(A)}$ onto the hyperboloid amounts to dropping the five dimensional radial component $\xi^{(A)R}$. Therefore, we continue to use the symbol $\xi^{(A)}$ to denote the four dimensional projected vector. Comparison of equation (7) with the conformal Killing equation $\nabla_a \xi_b + \nabla_b \xi_a = 2f g_{ab}$ shows that the four dimensional vectors $\xi^{(A)}$ are CKVs with corresponding conformal factors

$$f^{(A)} = -\frac{\kappa}{R}\xi_R^{(A)} \quad (8)$$

One can also derive a useful expression for the gradients of the conformal factors $f^{(A)}$. Along with equation (5), one has

$$0 = \nabla_a^{(5)} \xi_R^{(A)} \quad (9)$$

$$= \nabla_a \xi_R^{(A)} - \Gamma_{aR}^b \xi_b^{(A)}. \quad (10)$$

Plugging in $\Gamma_{aR}^b = \frac{1}{R}\delta_{ab}$ and using equation (8) then gives

$$\nabla_a f^{(A)} = -\kappa H^2 \xi_a^{(A)} \quad (11)$$

We see that the gradients of the conformal factors are simply proportional to the CKV's themselves.

Finally, we will need the explicit forms of the dS CKV's $\xi^{(A)}$ in the flat cosmological coordinates defined in (4),

$$\begin{aligned} \xi^{(0)} &= \left(\cosh(Ht) + \frac{r^2 H^2}{2} e^{Ht} \right) \left(\frac{\partial}{\partial t} \right) - H x^k e^{-Ht} \left(\frac{\partial}{\partial x^k} \right) \\ \xi^{(4)} &= \left(-\sinh(Ht) + \frac{H^2 r^2}{2} e^{Ht} \right) \left(\frac{\partial}{\partial t} \right) - H x^k e^{-Ht} \left(\frac{\partial}{\partial x^k} \right) \\ \xi^{(k)} &= -x^k e^{Ht} \left(\frac{\partial}{\partial t} \right) + \frac{1}{H} e^{-Ht} \left(\frac{\partial}{\partial x^k} \right) \end{aligned} \quad (12)$$

We note that the particular linear combination of CKV's, $\zeta = \xi^{(0)} - \xi^{(4)}$, is given simply by

$$\zeta = e^{Ht} \left(\frac{\partial}{\partial t} \right) \quad (13)$$

and is everywhere timelike and future directed. The CKV ζ generates translations in the conformal time parameter η defined by $d\eta = \exp(-Ht)dt$. We will see below that it is a conserved charge associated with ζ that satisfies a positivity relation.

Killing Spinors

We now turn to spinor fields in $(A)dS$. The Killing spinors of the flat embedding spacetimes lead to Killing spinors in $(A)dS$ that are constant with respect to the projection of the 5 dimensional covariant derivative operator onto the $(A)dS$ hyperboloid. Denoting this projected derivative operator by $\check{\nabla}_a$, we have

$$\check{\nabla}_a \eta = \nabla_a \eta + \frac{1}{2} \omega_a^{\hat{b}\hat{R}} \gamma_{\hat{b}\hat{R}} \eta, \quad (14)$$

where hats denote orthonormal frame indices. The necessary component of the spin connection for the metric (2) is given by $\omega_a^{\hat{b}\hat{R}} = \frac{\kappa}{R} e^{\hat{b}}_a$, where $e^{\hat{b}}_a$ are the components of the vierbein for the $(A)dS$ metric g_{ab} . This yields the expression

$$\check{\nabla}_a \eta = \nabla_a \eta + \frac{\kappa}{2R} \gamma_a \gamma_{\hat{5}} \eta, \quad (15)$$

Both dS and AdS inherit four complex solutions to $\check{\nabla}_a \phi = 0$, which we call Killing spinors, from the flat space embeddings. The explicit solutions for the Killing spinors of dS in the flat cosmological coordinates of (3) are given by

$$\eta(\vec{x}, t) = \left(1 - \frac{H}{2} x^k \gamma_k (\gamma^{\hat{5}} + \gamma^{\hat{t}})\right) \left(e^{\frac{Ht}{2}} \alpha_- + e^{-\frac{Ht}{2}} \alpha_+\right) \quad (16)$$

where α_{\pm} are constant spinors satisfying the conditions

$$(\gamma^{\hat{5}} \mp \gamma^{\hat{t}}) \alpha_{\pm} = 0. \quad (17)$$

Below, we will find it convenient to work with another form of the derivative operator. If one takes $\eta = \frac{1}{\sqrt{2}}(1 + c\gamma^{\hat{5}})\psi$ with $c^2 = -\kappa$, then one finds that

$$\check{\nabla}_a \eta = \frac{1}{\sqrt{2}}(1 + c\gamma^{\hat{5}})\hat{\nabla}_a \psi \quad (18)$$

where the new derivative operator $\hat{\nabla}_a$ is given by

$$\hat{\nabla}_a \psi = \nabla_a \psi + \frac{c}{2R} \gamma_a \psi. \quad (19)$$

With $c = 1$ for AdS , $\hat{\nabla}_a$ is the super-covariant derivative operator that comes from $N = 2$ gauged supergravity as given in *e.g.* [19]. With $c = i$ for dS , $\hat{\nabla}_a$ is the supercovariant derivative operator arising in the classical deSitter supergravity theory discussed in [12].

Bilinears in the Killing spinors give linear combinations of the (C)KV's for $(A)dS$ in the following way. If ψ satisfies $\hat{\nabla}_a \psi = 0$, one can show that the vector $\xi^a = -\bar{\psi} \gamma^a \psi$ satisfies

$$\nabla_a \xi_b = \begin{cases} \frac{i}{R} g_{ab} \bar{\psi} \psi, & dS \\ -\frac{1}{R} \bar{\psi} \gamma_{ab} \gamma_{\hat{5}} \psi, & AdS \end{cases} \quad (20)$$

We see that for AdS , the right hand side is anti-symmetric and ξ^a is therefore a Killing vector. For dS , the right hand side is proportional to the metric, and ξ^a is therefore a conformal Killing vector instead. It is also straightforward to show that the vectors $\xi^a = -\bar{\psi}\gamma^a\psi$ are everywhere future directed and timelike. Specializing to dS in the cosmological coordinates (3), one finds that the particular CKV ζ^a generating translations in the conformal time coordinate η is obtained by taking $\alpha_+ = 0$ in equation (16).

The situation is reversed if one looks at the vectors $\chi^a = -\bar{\psi}\gamma_5\gamma^a\psi$, which satisfy

$$\nabla_a\chi_b = \begin{cases} -\frac{i}{R}\bar{\psi}\gamma_{ab}\gamma_5\psi, & dS \\ \frac{1}{R}g_{ab}\bar{\psi}\psi, & AdS \end{cases} \quad (21)$$

The vectors χ^a are therefore Killing vectors for dS and conformal Killing vectors for AdS . Note that unlike the case of flat spacetime, the Killing vector one obtains from the Killing spinors are not irrotational, *i.e.* the antisymmetric piece $\nabla_{[a}\xi_{b]}$ is nonzero.

3 A Positive Energy Theorem for DeSitter

3.1 Derivation of the Spinor Identity

In this section, we derive a Gauss's law identity for spinor fields in 4-dimensional, asymptotically (A)dS spacetimes using the derivative operator $\hat{\nabla}_a$. Following the covariant approach of reference [20], we start by defining the Nester form $B^{ab} = \bar{\psi}\gamma^{abc}\hat{\nabla}_c\psi$. It is then straightforward to show that the divergence of twice the real part of B^{ab} can be written as

$$\begin{aligned} \nabla_a(B^{ab} + B^{ab*}) &= -(G^b_c + \Lambda g^b_c)\bar{\psi}\gamma^c\psi + 2(\overline{\hat{\nabla}_a\psi})\gamma^{abc}(\hat{\nabla}_c\psi) \\ &\quad + \frac{1}{R}(c - c^*)\left(\bar{\psi}\gamma^{ab}(\hat{\nabla}_a\psi) - \overline{(\hat{\nabla}_a\psi)}\gamma^{ab}\psi\right), \end{aligned} \quad (22)$$

where G_{ab} is the Einstein tensor and we have made use of the identity $\gamma^{abc}\nabla_a\nabla_c\psi = -\frac{1}{2}G^{bc}\gamma_c\psi$. Recalling that $c = i$ for dS and $c = 1$ for AdS , we see that the coefficient of last term in equation (22) vanishes for AdS .

Now choose a spacelike slice Σ with timelike unit normal n^a . The spacetime metric can then be written as $g_{ab} = -n_a n_b + q_{ab}$, with q_{ab} the metric on Σ . Stokes theorem implies that a 2-form field A_{ab} satisfies

$$\int_V (\nabla_b A^{bc})n_c = -\frac{1}{2} \int_{\partial V} ds_{ab} A^{ab}, \quad (23)$$

where the spatial volume V lies within Σ and has boundary ∂V . For a spinor field ψ define the quantity Q_ψ by the surface integral

$$Q_\psi = -\frac{1}{16\pi} \int_{\partial V} ds_{ab} (B^{ab} + B^{ab*}) \quad (24)$$

$$= -\frac{1}{16\pi} \int_{\partial V} ds_{ab} \left(\bar{\psi}\gamma^{abc}(\hat{\nabla}_c\psi) - \overline{(\hat{\nabla}_c\psi)}\gamma^{abc}\psi\right) \quad (25)$$

Stokes theorem together with the identity (22) then imply that Q_ψ may be alternatively written as a volume integral

$$Q_\psi = \frac{1}{8\pi} \int_V \left[(G_{ab} + \Lambda g_{ab}) \xi^a n^b + q^{ij} (\hat{\nabla}_i \psi^\dagger) \hat{\nabla}_j \psi - (\gamma^i \hat{\nabla}_i \psi)^\dagger (\gamma^j \hat{\nabla}_j \psi) \right] \quad (26)$$

$$- \frac{1}{R} (c - c^*) \left(\psi^\dagger (\gamma^j \hat{\nabla}_j \psi) - (\gamma^j \hat{\nabla}_j \psi)^\dagger \psi \right) \Big], \quad (27)$$

where $\xi^a = -\bar{\psi} \gamma^a \psi$ and the indices i, j are raised and lowered with the spatial metric q_{ij} . If we now impose that the spinor field satisfy the Dirac-Witten equation

$$\gamma^j \hat{\nabla}_j \psi = 0, \quad (28)$$

then on solutions to the Einstein equations $G_{ab} + \Lambda g_{ab} = 8\pi T_{ab}$ one has the relation

$$Q_\psi = \frac{1}{8\pi} \int_\Sigma \left(8\pi T_{ab} \xi^a n^b + q^{ij} (\hat{\nabla}_i \psi)^\dagger \hat{\nabla}_j \psi \right) \quad (29)$$

The second term is manifestly non-negative, vanishing only if $\hat{\nabla}_i \psi = 0$. Since the vector field $\xi^a = -\bar{\psi} \gamma^a \psi$ is by construction everywhere future directed and timelike, the first term will also be non-negative provided that the stress-energy tensor satisfies the dominant energy condition (see *e.g.* [21]). Subject to this condition on the stress-energy tensor, we have then proven that

$$Q_\psi \geq 0. \quad (30)$$

For anti-deSitter spacetimes this is the standard result of [5]. The charge Q_ψ in this case is a linear combination of the conserved charges defined in reference [6], corresponding to the Killing vectors of anti-deSitter spacetimes. For deSitter spacetimes, equation (30) is a new result. The nature of the charge Q_ψ in deSitter remains to be explored. Since $\xi^a = -\bar{\psi} \gamma^a \psi$ is a conformal Killing vector in this case, we expect that the charge Q_ψ is a conserved quantity related to the conformal Killing vectors of deSitter spacetimes. Indeed, since B^{ab} is antisymmetric, the divergence of left hand side of (22) is identically equal to zero,

$$\nabla_b J^b \equiv \nabla_b \nabla_a (B^{ab} + B^{ab*}) = 0 \quad (31)$$

Hence Q_ψ is the volume integral of the time component of a conserved current. We further discuss the interpretation of Q_ψ in section (5).

The above analysis can be repeated using the more general two form $A^{ab} = \bar{\psi} (\gamma^5)^p \gamma^{abc} \hat{\nabla}_c \psi$. For $p = 0$ this is the same as the case we have been analyzing, and Q_ψ depends on the CKV's of deSitter. For $p = 1$, Q_ψ depends on the KV's of deSitter. This construction is summarized in the Appendix. However, in the case of the KV's we show that the $\hat{\nabla} \psi$ contribution to the volume integrand is not positive definite. Therefore even if there are no matter sources, $T_{ab} = 0$, the charge generated by a Killing vector does not appear to be positive definite⁶. Alternatively, one could consider the KV case for perturbations off deSitter. However, none of the KV's is timelike everywhere, so again the integrand is not positive definite.

⁶See however the argument in reference [16] that the Abbott & Deser mass is nonetheless positive semi-definite.

3.2 Evaluation of Q_ψ in the Asymptotic Limit

We now want to rewrite the boundary integral expression for the charge Q_ψ (25) in terms of the asymptotic behavior of the metric. We begin by writing the spacetime metric in the asymptotic region as $g_{ab} = \tilde{g}_{ab} + h_{ab}$. Here \tilde{g}_{ab} is the deSitter metric⁷. We will work with flat cosmological coordinates, in which the asymptotic deSitter metric takes the form (3). The rate of falloff we require for the metric perturbation h_{ab} in the asymptotic region is discussed below. We assume that in the asymptotic region, the spinor field has the form $\psi = \tilde{\psi} + \delta\psi$, where $\tilde{\psi}$ is a deSitter Killing spinor and $\delta\psi$ is falling off to zero.

We are seeking to write the boundary integral expression for Q_ψ in terms of h_{ab} . The most conceptually straightforward way to do this would be to start with a general metric perturbation h_{ab} satisfying the asymptotic conditions, solve the Dirac-Witten equation subject to the boundary condition that ψ approach a Killing spinor, and then substitute into the expression (25) for Q_ψ . In the asymptotically flat case [2], this calculation leads to the result $Q_\psi = -P_a \xi^a$, where ξ^a is the future directed timelike Killing vector built out of the asymptotic Killing spinor. Positivity of Q_ψ then implies that the ADM 4-momentum P^a is future directed, timelike or null.

In the present case, this approach proves to be quite cumbersome due to the fact that the extrinsic curvature of the background is nonzero. Instead we will follow the method used in reference [22]. We will also restrict our analysis to spinor fields ψ that approach Killing spinors which are eigenfunctions of $\gamma^{\hat{t}}$. These have $\alpha_+ = 0$ in equation (16). As noted above, the vector $\xi^a = -\bar{\psi}\gamma^a\psi$ formed from these spinors is the conformal time translation CKV. The main result of this section is the expression (50) for the charge Q_ψ corresponding to these spinors. The derivation of (50) is somewhat lengthy and the reader may want to simply skip ahead to this equation to focus on the results.

Following reference [22], we begin by splitting the derivative operator $\hat{\nabla}_a$ into a background piece and a perturbation $\hat{\nabla}_a = \hat{\tilde{\nabla}}_a + (\hat{\nabla}_a - \hat{\tilde{\nabla}}_a)$ and writing the Nester form B^{ab} as

$$B^{ab}[\hat{\nabla}] = B^{ab}[\hat{\tilde{\nabla}}] + B^{ab}[\hat{\nabla} - \hat{\tilde{\nabla}}]. \quad (32)$$

We first show that the background part of the boundary integral vanishes,

$$\int_{\partial\Sigma} ds_{ab} B^{ab}[\hat{\tilde{\nabla}}] = 0. \quad (33)$$

As indicated above, we take the Killing spinor $\tilde{\psi}$ to be given by $\tilde{\psi} = \frac{1}{\sqrt{2}}(1 - i\gamma^{\hat{5}})e^{Ht/2}\alpha_+$. With this choice the boundary integral (33) is found to depend only on the $1/r$ terms in $\delta\psi$. This choice of $\tilde{\psi}$ is an eigenspinor⁸ of $\gamma^{\hat{t}}$, satisfying $\gamma^{\hat{t}}\tilde{\psi} = -i\tilde{\psi}$. The argument now proceeds by the following steps.

⁷In the following sections, tilde's will denote deSitter quantities. For example, the Dirac matrices $\tilde{\gamma}^a$ satisfy the algebra $\{\tilde{\gamma}^a, \tilde{\gamma}^b\} = 2\tilde{g}^{ab}$.

⁸The Killing spinor constructed from α_+ is not an eigenspinor of $\gamma^{\hat{t}}$ and hence the argument below would not imply finiteness of Q_ψ in this case.

- 1: Show that the boundary integral vanishes if $\delta\psi$ also satisfies $\gamma^{\hat{t}}\delta\psi = -i\delta\psi$.
- 2: Determine consistent falloff conditions on the metric perturbation h_{ab} , such that the Einstein constraint equations are satisfied.
- 3: Show that to order $1/r$, $\delta\psi$ satisfies the requirement of step 1.

Step 1: We want to show that equation (33) holds, if $\gamma^{\hat{t}}\delta\psi = -i\delta\psi$. To leading order in $1/r$ we have,

$$B^{ab}(\hat{\nabla}) = \tilde{\psi}\gamma^{abc}(\tilde{\nabla}_c + i\frac{H}{2}\gamma_c)\delta\psi \quad (34)$$

Integration by parts together with the relation $\hat{\nabla}_a\tilde{\psi} = 0$ then yields

$$\int_{\partial\Sigma} ds_{ab}\tilde{\psi}\gamma^{abc}\tilde{\nabla}_c\delta\psi = i\frac{H}{2}\int_{\partial\Sigma} ds_{ab}\tilde{\psi}\gamma_c\gamma^{abc}\delta\psi. \quad (35)$$

Using this result, the surface integral of $B^{ab}[\hat{\nabla}]$ in (33) can then be written as $2iH\int_{\partial\Sigma} ds_{tj}\tilde{\psi}\gamma^t\gamma^j\delta\psi$. The integrand then vanishes if both $\tilde{\psi}$ and $\delta\psi$ are eigenspinors of $\gamma^{\hat{t}}$ with the same eigenvalue. The rest of the argument below shows that $\delta\psi$ is indeed such an eigenspinor to order $1/r$.

Step 2: We must now discuss the falloff conditions on h_{ab} in the asymptotic region. We require that the rate of fall off is general enough to include solutions to the Einstein constraint equations corresponding to localized stress energy sources with non-zero monopole moments. That is, we want the resulting positive mass theorem to include the case of a galaxy in deSitter⁹.

Schwarzschild-deSitter spacetime is useful as a reference, since we can solve the Dirac-Witten equation in this case. In static coordinates the metric is

$$ds^2 = -F(R)dT^2 + \frac{1}{F(R)}dR^2 + R^2d\Omega^2, \quad F(R) = 1 - \frac{2M}{R} - \frac{\Lambda}{3}R^2. \quad (36)$$

The Killing vector $\partial/\partial T$ is, of course, only timelike in the region between the deSitter and black hole horizons. Transforming to cosmological coordinates, the metric (36) becomes the McVittie metric [23]

$$ds^2 = -\frac{(1 - \frac{M}{2ar})^2}{(1 + \frac{M}{2ar})^2} dt^2 + a^2(1 + \frac{M}{2ar})^4 \delta_{ij}dx^i dx^j, \quad a(t) = e^{Ht} \quad (37)$$

Then the equation $\gamma^j\hat{\nabla}_j\psi = 0$ is solved exactly for the McVittie metric by the spinors

$$\psi = \Omega^{-2}\tilde{\psi}, \quad \Omega = 1 + \frac{M}{2ar} \quad (38)$$

where $\tilde{\psi}$ is a Killing spinor of the background deSitter spacetime. Substituting ψ into the boundary term, one explicitly calculates $Q_\psi = a(t)M$. Both the metric and

⁹See also Shiromizu et al [16] for an alternative discussion of falloff conditions in asymptotically deSitter spacetimes, which arrives at the same conditions.

the spinor fields fall off like r^{-1} , as one expects. However, note that simple power counting suggests that the boundary integral (35) diverges: $\tilde{\psi}$ is constant, $\delta\psi$ goes like r^{-1} , and the area element goes like r^2 . The fact that (35) is zero (rather than infinity) is due to the orthogonality properties of the spinors involved, which follow from the fact that ψ is an eigenfunction of $\gamma^{\hat{t}}$.

Turning to the general case, let Σ be a spatial slice with unit normal n^a , spatial metric q_{ij} , extrinsic curvature K^{ij} , and let $\rho = T_{ab}n^an^b$ be the matter density. To determine appropriate falloff conditions we use the Hamiltonian constraint equation

$$R^{(3)} + K^2 - K^{ij}K_{ij} - 2\Lambda = 16\pi\rho \quad (39)$$

If ρ is a compact source, then in the far field the spatial metric q_{ij} satisfies a Poisson-type equation, and hence the perturbation to q_{ij} vanishes like $1/r$. Therefore we require that in the asymptotic region, the spatial metric has the form

$$q_{ij} = \tilde{q}_{ij} + \mathcal{O}(1/r). \quad (40)$$

Let $k^{ij} = K^{ij} - \tilde{K}^{ij}$ be the perturbation to the extrinsic curvature. In the case of the McVittie metric (37), the extrinsic curvature is given by $K^{ij} = Hq^{ij}$, which exactly balances the cosmological constant term in the Hamiltonian constraint equation (39). In general, we find it judicious to let $K^{ij} = Hq^{ij} + l^{ij}$, so that the perturbation to the extrinsic curvature becomes $k^{ij} = H(q^{ij} - \tilde{q}^{ij}) + l^{ij}$. Equation (40) implies that the 3-dimensional scalar curvature at infinity vanishes as $R^{(3)} \sim \mathcal{O}(1/r^3)$. Substituting into the constraint equation then implies that

$$l^{ij} \sim \mathcal{O}(1/r^3) \quad (41)$$

The order $\mathcal{O}(1/r)$ corrections to the extrinsic curvature enter only from the Hq^{ij} term. The perturbation l^{ij} must fall off more rapidly because the background extrinsic curvature is nonzero¹⁰. Note that the trace of the extrinsic curvature is given by $K = 3H + l$. The Einstein momentum constraint does not add any new information.

Step 3: We now show that $\gamma^{\hat{t}}\delta\psi = -i\delta\psi$ to order $1/r$. The spinor field ψ must solve the Dirac-Witten equation $\gamma^k\hat{\nabla}_k\psi = 0$. In detail this is

$$\gamma^k\partial_k\psi + \frac{1}{4}\omega_k^{\hat{a}\hat{b}}\gamma^k\gamma_{\hat{a}\hat{b}}\psi + \frac{3}{2}iH\psi = 0, \quad (42)$$

where $\omega_c^{\hat{a}\hat{b}} = e^{\hat{a}d}(\partial_c e_d^{\hat{b}} - \Gamma_{cd}^e e_e^{\hat{b}})$ is the spin connection. We choose coordinates that $h_{tj} = 0$. This means that the normal to Σ is asymptotically in the direction of $\frac{\partial}{\partial t}$. One then finds $\omega_k^{\hat{t}\hat{i}} = K^i_k$ and hence

$$\gamma^k\omega_k^{\hat{t}\hat{i}}\gamma_{\hat{i}\hat{i}} = 3H + l. \quad (43)$$

¹⁰In n spacetime dimensions the fall off conditions become $\delta q_{ij} \sim r^{-n+3}$, $l^{ij} \sim r^{-n+1}$.

The ω_i^{jk} components of the spin connection vanish in the background. The Dirac-Witten equation for ψ becomes

$$\gamma^k \partial_k \psi + \frac{1}{4} \omega_k^{ij} \gamma^k \gamma_{ij} \psi + \frac{3}{2} i H (1 - i \gamma^{\hat{t}}) \psi + \frac{1}{2} l \psi = 0 \quad (44)$$

We want to show that (44) has solutions which are eigenspinors of $\gamma^{\hat{t}}$, at least through the first correction $\delta\psi$. If this is true, then the third term in equation (44) vanishes. Assuming that this is the case, now multiply equation (44) by $\gamma^{\hat{t}}$ and rewrite it as a differential equation for $\gamma^{\hat{t}}\psi$. The first two terms change sign, but the last one does not. However, $l \sim 1/r^3$ while $\omega_k^{\hat{i}\hat{j}} \sim 1/r^2$. The l term is then of higher order.

Therefore to leading order, we want a solution to

$$\gamma^k \partial_k \delta\psi = -\frac{1}{4} \omega_k^{ij} \gamma^k \gamma_{ij} \tilde{\psi}, \quad (45)$$

such that $\delta\psi$ satisfies $\gamma^{\hat{t}}\delta\psi = -i\delta\psi$. Let N_A be the eigenspinors of $\gamma^{\hat{t}}$ with eigenvalue $-i$, and P_A be the eigenspinors with eigenvalue $+i$, with $A = 1, 2$ in four dimensions. Choose the indexing such that $\tilde{\psi} = \exp(Ht)N_1$. Multiplying by a spatial Dirac matrix flips the eigenvalue of $\gamma^{\hat{t}}$, so that that $(\gamma^j N_A) = C_{AB}^j P_B$, where C_{AB}^j is a constant matrix. Now expand the perturbation $\delta\psi$ to the spinor as $\delta\psi = F_B N_B$, where the F_B are unknown functions to be solved for. In the differential equation (45) for $\delta\psi$, we have

$$\omega_k^{ij} \gamma^k \gamma_{ij} = -\Gamma_{kl}^n \gamma^k \gamma^l \gamma_n = \alpha_i \gamma^i, \quad (46)$$

where the α_i are known functions in terms of derivatives of h_{ij} . Substituting this into equation (45) we then have

$$[C_{BC}^k \partial_k F_B + e^{Ht} \alpha_i C_{1C}^i] N_C = 0. \quad (47)$$

This is a set of two first order PDEs for the two unknown functions F_B , and therefore one expects that generically there is a solution.

This completes the argument that the integral $\int_{\partial\Sigma} ds_{ab} B^{ab} [\hat{\nabla}^{dS}]$ vanishes. Note that at next order in powers of $1/r$ the l term does contribute to the differential equation (44), and therefore it is not possible to find solutions which are eigenspinors of $\gamma^{\hat{t}}$ everywhere in the volume Σ . However, things work out just right to have the $B^{ab}(\hat{\nabla})$ contribution vanish at infinity. Equation (24) for the charge Q_ψ now reduces to

$$Q_\psi = -\frac{1}{16\pi} \int ds_{ab} \left(B^{ab} [\hat{\nabla} - \hat{\tilde{\nabla}}] + B^{ab*} [\hat{\nabla} - \hat{\tilde{\nabla}}] \right) \quad (48)$$

The difference in derivative operators acting on the spinor field ψ can then be written as

$$(\nabla_a - \tilde{\nabla}_a) \psi = \frac{1}{4} g_{bc} C_{da}^c \gamma^{bd} \psi, \quad (49)$$

with $C_{bc}^a = g^{ad} (\tilde{\nabla}_b h_{cd} + \tilde{\nabla}_c h_{bd} - \tilde{\nabla}_d h_{bc})$ and $h_{ab} = g_{ab} - \tilde{g}_{ab}$. Substitution then yields an explicit expression for the charge Q_ψ

$$Q_\psi = \frac{1}{16\pi} \int ds_{ab} [\xi^a (\tilde{\nabla}_d h^{db} - \tilde{\nabla}^b h) - \xi^b (\tilde{\nabla}_d h^{da} - \tilde{\nabla}^a h) + \xi^c (\tilde{\nabla}^b h_c^a - \tilde{\nabla}^a h_c^b)] \quad (50)$$

where as before $\xi^a = -\tilde{\psi}\gamma^a\tilde{\psi}$ and $h = \tilde{g}^{ab}h_{ab}$. This is a familiar expression from the asymptotically flat case. If $\tilde{\nabla}_a$ were instead the flat spacetime covariant derivative operator and ξ^a a time translation Killing vector, then equation (50) would be the expression for the ADM mass.

4 CKV's and Conserved Charges

4.1 General Construction

In section (3) we constructed charges Q_ψ associated with the conformal Killing vectors of the background deSitter spacetime. These charges are conserved. In particular for the conformal time translation CKV we showed that $Q_\psi \geq 0$. This result raises the question, what is the physical significance of these charges? One usually associates conserved charges with symmetries. However, this construction involves the conformal symmetries. In this section, we will see that for asymptotically (anti-)deSitter spacetimes, there exist conserved charges corresponding to the background CKV's as well as those corresponding to KV's.

Our result follows from an extension of the methods of reference [6]. The Abbott and Deser construction begins with spacetimes that are asymptotic at infinity to a fixed background spacetime, which has some number of Killing vectors. For each background Killing vector, they showed that there exists a conserved current and further that the corresponding conserved charge, obtained by integrating the time component of the current over a spatial slice, can be re-expressed as a boundary integral at spatial infinity. The construction in [6] is covariant in nature. We note that these results may also be obtained using the Hamiltonian formalism for general relativity, as a special case of the integral constraint vector (ICV) construction of reference [24] (see also [25]). Here we will follow the covariant approach.

Consider a spacetime with metric g_{ab} and stress energy tensor T_{ab} that satisfies the Einstein equations with cosmological constant Λ ,

$$R_{ab} - \frac{1}{2}g_{ab}R + \Lambda g_{ab} = 8\pi T_{ab}. \quad (51)$$

Assume that the spacetime is asymptotic at spatial infinity to a fixed background spacetime with metric \tilde{g}_{ab} and vanishing stress energy¹¹ that also solves the Einstein equations with cosmological constant Λ . Define the tensor h_{ab} to be the difference between the spacetime metric and the fixed background metric.

$$h_{ab} = g_{ab} - \tilde{g}_{ab}. \quad (52)$$

Note that h_{ab} is not assumed to be small in the interior of the spacetime, but must vanish at an appropriate rate near spatial infinity (see Section 3.2).

¹¹The Hamiltonian approach of [24] allows for background stress-energy tensor $\tilde{T}_{ab} \neq 0$. This involves introducing vector fields more general than KV's and CKV's, which are the ICV's mentioned above.

The Abbott and Deser construction [6] proceeds by expanding the curvature of g_{ab} in powers of h_{ab} . The Einstein equations (51) are rewritten keeping terms linear in h_{ab} on the left hand side and moving all the nonlinear terms $G_{ab}^{(NL)}$ to the right hand side of the equation, giving

$$R_{ab}^{(L)} - \frac{1}{2}\tilde{g}_{ab}R^{(L)} - \Lambda h_{ab} = 8\pi T_{ab} - G_{ab}^{(NL)} \quad (53)$$

where $R_{ab}^{(L)}$ and $R^{(L)}$ are the terms linear in h_{ab} in the expansions for the Ricci curvature and scalar curvature of g_{ab} . Collecting all the nonlinear terms and the matter stress-energy into the quantity $\mathcal{T}_{ab} = 8\pi T_{ab} - G_{ab}^{(NL)}$, equation (53) can now be processed into the useful form

$$\mathcal{T}^{ac} = \tilde{\nabla}_b \tilde{\nabla}_d K^{abcd} + \frac{1}{2}\tilde{R}^c{}_{bde} K^{abde}, \quad (54)$$

where indices have been raised using the background metric and $\tilde{\nabla}_a$ is the covariant derivative operator for the background metric \tilde{g}_{ab} . The tensor K^{abcd} is defined by

$$K^{abcd} = \frac{1}{2} [\tilde{g}^{ad} H^{bc} + \tilde{g}^{bc} H^{ad} - \tilde{g}^{ac} H^{bd} - \tilde{g}^{bd} H^{ac}], \quad (55)$$

where $H^{ab} = h^{ab} - \frac{1}{2}\tilde{g}^{ab}h$, and has the same symmetries as the Riemann tensor.

The Bianchi identity for the Einstein tensor implies that $\tilde{\nabla}_a \mathcal{T}^{ab} = 0$. It then follows that if ξ^a is a Killing vector of the background metric, then the current

$$\mathcal{J}^a = \mathcal{T}^{ab}\xi_b \quad (56)$$

is conserved with respect to the background derivative operator,

$$\tilde{\nabla}_a \mathcal{J}^a = 0. \quad (57)$$

A conserved charge Q_ξ is now obtained by integrating the normal component of the current over a spatial slice Σ with respect to the background volume element,

$$Q_\xi = \int_\Sigma d^3x \sqrt{-\tilde{g}} \mathcal{J}^a n_a. \quad (58)$$

Abbott and Deser [6] then further show that the charge Q_ξ can be re-expressed as an integral over the boundary $\partial\Sigma$ of the spatial slice Σ at spatial infinity.

We now show how to extend these results to include the CKVs of background (anti-)deSitter spacetimes. In fact we will derive a considerably more general result. Take any vector field ξ^a , not necessarily a background KV or CKV, and contract ξ^a with both sides of equation (54). After some algebra it follows that

$$\begin{aligned} \mathcal{T}^{ab}\xi_b &= \tilde{\nabla}_c \left((\tilde{\nabla}_d K^{abcd})\xi_b - K^{abdc}\tilde{\nabla}_d \xi_b \right) \\ &+ \left(K^{abcd}\tilde{\nabla}_d \tilde{\nabla}_c + \frac{1}{2}K^{aecd}\tilde{R}^b{}_{ecd} \right) \xi_b \end{aligned} \quad (59)$$

By judiciously moving terms from the right hand side to the left, equation (59) can be put in a Gauss's law form

$$\mathcal{C}^a = \tilde{\nabla}_b \mathcal{B}^{ab}, \quad (60)$$

where \mathcal{B}^{ab} and \mathcal{C}^a are given by

$$\mathcal{B}^{ab} = (\tilde{\nabla}_c K^{abdc}) \xi_d - K^{abdc} \tilde{\nabla}_{[c} \xi_{d]} \quad (61)$$

$$\mathcal{C}^a = \mathcal{T}^{ab} \xi_b + \tilde{\nabla}_c (K^{adbc} \tilde{\nabla}_{(d} \xi_{b)}) + \left(K^{acbd} \tilde{\nabla}_d \tilde{\nabla}_c + \frac{1}{2} K^{aecd} \tilde{R}^b{}_{ecd} \right) \xi_b, \quad (62)$$

The symmetries of K^{abcd} imply that \mathcal{B}^{ab} is an antisymmetric tensor. Therefore (60) implies that the vector field \mathcal{C}^a is divergenceless with respect to the background derivative operator, $\tilde{\nabla}_a \mathcal{C}^a = 0$. If ξ^a is a background Killing vector, then the second and third terms in \mathcal{C}^a vanish, and \mathcal{C}^a reduces to the current \mathcal{J}^a defined above. Equation (60) is then the result of Abbott and Deser [6]. More generally, however, the Gauss's law identity (60) holds, and hence also the conservation law $\tilde{\nabla}_a \mathcal{C}^a = 0$, holds for any vector field ξ^a . This result may seem surprising. However, we should recall that Wald [14] has shown that any vector field, acting as a generator of diffeomorphisms, gives rise to a conserved current and corresponding Noether charge. We expect that our result has an interpretation in this context.

Taking a spatial surface Σ with unit timelike normal n^a and boundary $\partial\Sigma$, define the quantity Q_ξ by the surface integral

$$Q_\xi = \frac{1}{8\pi} \int_{\partial\Sigma} ds_{ab} \mathcal{B}^{ab}. \quad (63)$$

Stokes theorem and equation (60) then imply that Q_ξ can also be written as the volume integral

$$Q_\xi = \frac{1}{8\pi} \int_\Sigma \sqrt{\tilde{g}} \mathcal{C}^a n_a. \quad (64)$$

The form of the boundary integral (63) does not depend on the presence, or absence, of a cosmological constant and it therefore holds, in particular for asymptotically flat spacetimes. In this case, plugging in the translational Killing vectors of Minkowski spacetime for ξ^a yields the usual expressions for the components of the ADM 4-momentum. Turning to asymptotically deSitter spacetimes, the expression for the Abbott & Deser mass [6] is obtained by inserting the time translation Killing vector of deSitter spacetime in static coordinates.

In Schwarzschild-deSitter, the charge for the CKV $\zeta = a(t) \frac{\partial}{\partial t}$ is $Q_\zeta = a(t)M$. One may wonder about the fact that Q_ζ depends on time t , since we have shown above that it is a conserved charge. The time dependence comes about because there is a nonzero flux of the spatial components of the current \mathcal{C}^k through spatial infinity. One can explicitly compute the flux of \mathcal{C}^k for Schwarzschild-deSitter, and verify that it equals $\partial_t Q_\zeta = \dot{a}M$.

4.2 Equivalence of Q_ψ and Q_ξ for CKV's

Now consider asymptotically deSitter spacetimes, and let ξ^a be one of the conformal Killing vectors of the background deSitter spacetime. The main result of this section is to demonstrate that the charge Q_ξ , in this case, is related to the charge Q_ψ constructed for deSitter CKV's in section (3). We will check this both for the surface integral and the volume integral expressions for the two charges agree. For the surface integrals, this is straightforward. Plugging a deSitter CKV into the boundary integral expression for Q_ξ (63), the resulting expression can easily be put in the form of the boundary integral for Q_ψ given in equation (50) with the identification $\xi^a = -\tilde{\psi}\tilde{\gamma}^a\tilde{\psi}$. In particular, the term proportional to the shear $\tilde{\nabla}_{[b}\xi_{d]}$ of the CKV vanishes by virtue of equation (20).

Having checked that the boundary integral expressions for Q_ξ and Q_ψ agree, it then follows that the volume integral expressions must also agree. However, it is interesting to check this explicitly, in order to gain some intuition into the conserved charges associated with CKV's. Of course, it is not possible (or at least not plausible) for a general spacetime to demonstrate directly that the volume integrands for Q_ξ and Q_ψ agree to all orders in h_{ab} . We will work only to linear order in h_{ab} .

Even at linear order h_{ab} , checking agreement of the volume integrals for Q_ξ and Q_ψ is a more complicated task than checking agreement of the surface integral expressions. To see why, note that the charges Q_ξ and Q_ψ , regarded as surface integrals, depend respectively only on the behavior of the vector field ξ^a and the spinor field ψ in the asymptotic region. There are then infinitely many ways extend ξ and ψ in the interior that keep the corresponding charges fixed. These are essentially gauge degrees of freedom. In the present case, we have specified that ξ^a will be taken to be a deSitter CKV throughout the interior. However, we have not yet specified how to extend the spinor field ψ into the interior. There are two natural choices. One choice is to simply take ψ to be one of the deSitter Killing spinors $\tilde{\psi}$ throughout the spacetime. Then the boundary expression for Q_ψ is still given by (50). This has the disadvantage that the spinor will not satisfy the Dirac-Witten equation (28) and hence the positivity of Q_ψ is not manifest with this choice. However, this choice has the advantage that ξ^a and ψ are related by $\xi^a = -\tilde{\psi}\tilde{\gamma}^a\tilde{\psi}$ everywhere in the spacetime, not just in the asymptotic region. We will see that with this choice the linearized integrands of the volume integral expressions for the charges Q_ξ and Q_ψ agree at each point. A second natural choice for the extension of ψ is to impose the Dirac-Witten equation (28). This choice makes the volume integrand manifestly positive. We will see, in this case, that the volume integrals for Q_ξ and Q_ψ again match, as they must. However, the volume integrands do not match on a point by point basis.

The volume integrand for Q_ξ is simply the normal component of the vector \mathcal{C}^a . Let us start by specializing equation (62) for \mathcal{C}^a to the case of a deSitter CKV. Given the properties of deSitter CKV's $\tilde{\nabla}_a\xi_b = f\tilde{g}_{ab}$, with $\tilde{\nabla}_af = -\frac{\Lambda}{3}\xi_a$, it follows that

$$\tilde{\nabla}_b\tilde{\nabla}_c\xi_d = -\tilde{R}_{cdb}{}^e\xi_e - \frac{\Lambda}{3}(\tilde{g}_{cd}\xi_b + \tilde{g}_{bd}\xi_c - \tilde{g}_{bc}\xi_d). \quad (65)$$

Plugging this into equation (62) for \mathcal{C}^a then gives

$$\mathcal{C}^a = \mathcal{T}^{ab}\xi_b - \frac{\Lambda}{3}K^{a\alpha b\beta}\tilde{g}_{\alpha\beta}\xi_b + f\left(\tilde{\nabla}_\beta K^{a\alpha b\beta}\right)\tilde{g}_{\alpha b}, \quad (66)$$

which includes additional terms linear in h_{ab} , relative to the current $\mathcal{J}^a = \mathcal{T}^{ab}\xi_b$ in the Killing vector construction.

We now follow the first approach and take ψ to be a Killing spinor, satisfying

$$\hat{\nabla}_a\psi = 0, \quad (67)$$

everywhere in the interior. From section (3), the volume integrand for Q_ψ is the normal component of the conserved current \mathcal{D}^b defined by the right hand side of equation (22)

$$\begin{aligned} \mathcal{D}^b = & \left(G^b_c + \Lambda g^b_c\right)(-\bar{\psi}\gamma^c\psi) + 2\overline{(\hat{\nabla}_a\psi)}\gamma^{abc}(\hat{\nabla}_c\psi) \\ & + 2iH\left(\bar{\psi}\gamma^{ab}(\hat{\nabla}_a\psi) - \overline{(\hat{\nabla}_a\psi)}\gamma^{ab}\psi\right), \end{aligned} \quad (68)$$

The task is now to compare the currents \mathcal{C}^a and \mathcal{D}^a to linear order in h_{ab} , with the identification $\xi^a = -\bar{\psi}\tilde{\gamma}^a\tilde{\psi}$, where $\tilde{\gamma}^a$ are background gamma matrices satisfying $\{\tilde{\gamma}^a, \tilde{\gamma}^b\} = 2\tilde{g}^{ab}$. Making use of Einstein's equation, the first terms in equations (66) and (68) clearly agree to this order. To compare the other terms we need an expression for $\hat{\nabla}_a\psi$ to linear order, which is found to be

$$\hat{\nabla}_a\psi = -\frac{1}{4}(\tilde{\nabla}_b h_{ac})\tilde{\gamma}^{bc}\psi + \frac{i}{4}Hh_{ab}\tilde{\gamma}^b\psi + \mathcal{O}(h^2). \quad (69)$$

The second term in equation (68) is second order in $\hat{\nabla}_a\psi$ and therefore does not contribute to linear order in h_{ab} . Dirac matrix algebra reduces the third term in (68) to

$$2iH\left(\bar{\psi}\gamma^{ab}(\hat{\nabla}_a\psi) - \overline{(\hat{\nabla}_a\psi)}\gamma^{ab}\psi\right) = \frac{\Lambda}{3}(h^b_c - \tilde{g}^b_c h)(-\bar{\psi}\tilde{\gamma}^c\psi) - iH(\tilde{\nabla}^b h - \tilde{\nabla}_a h^{ab})\bar{\psi}\psi. \quad (70)$$

Making use of the relation for the conformal factor $f = iH\bar{\psi}\psi$ from equation (20) and the definition of K^{abcd} in equation (55) then shows the equality of (70) with the last two terms in equation (66) for \mathcal{C}^a . Therefore, we have shown that the volume integrands for the charges Q_ξ and Q_ψ agree pointwise to linear order,

$$Q_\psi = Q_\xi = \int_\Sigma \left(T^{ab}n^a\xi_b - \frac{\Lambda}{3}K^{acbd}\tilde{g}_{cd}n^a\xi_b + f\left(\tilde{\nabla}_d K^{acbd}\right)\tilde{g}_{cb}n^a\right) \quad (71)$$

$$= \int_\Sigma \left(T_{ab}(-\bar{\psi}\tilde{\gamma}^a\tilde{\psi})n^b - \frac{iH}{4\pi}\left(\tilde{\psi}^\dagger(\gamma^j\hat{\nabla}_j\tilde{\psi}) - (\gamma^j\hat{\nabla}_j\tilde{\psi})^\dagger\tilde{\psi}\right)\right). \quad (72)$$

We now turn to the second alternative, extending ψ to the interior by imposing the Dirac-Witten condition $\gamma^a\hat{\nabla}_a\psi = 0$ everywhere. The spinor field ψ then has a

first order correction, $\psi = \tilde{\psi} + \delta\psi$. The boundary term is unchanged. With the Dirac-Witten condition imposed, the volume expression for Q_ψ reduces to the manifestly positive expression in equation (29). Further, because $\hat{\nabla}_i\psi$ is itself first order in h_{ab} , to linear order Q_ψ given by the simple expression

$$Q_\psi = \int_\Sigma T_{ab}(-\tilde{\psi}\tilde{\gamma}^a\tilde{\psi})n^b \quad (73)$$

Consistency between equations (73) and (72) requires that

$$\int_\Sigma \left(\tilde{\psi}^\dagger(\gamma^j\hat{\nabla}_j\tilde{\psi}) - (\gamma^j\hat{\nabla}_j\tilde{\psi})^\dagger\tilde{\psi} \right) = 0. \quad (74)$$

To verify this, use $\tilde{\psi} = \psi - \delta\psi$. Then to linear order in h_{ab} $\tilde{\psi}^\dagger(\gamma^j\hat{\nabla}_j\tilde{\psi}) = \tilde{\nabla}_j(\tilde{\psi}^\dagger\gamma^j\delta\psi)$. The integrand in equation (74) is thus a total divergence and can be rewritten using Stokes theorem as

$$\frac{1}{8\pi} \int_{\partial\Sigma} da_j(-2iH) \left(\tilde{\psi}^\dagger\gamma^j\delta\psi - \delta\psi^\dagger\gamma^j\tilde{\psi} \right). \quad (75)$$

Our earlier result from Section 3.2 $\gamma^{\hat{t}}\delta\psi = -i\delta\psi$ then implies that $\tilde{\psi}^\dagger\gamma^j\delta\psi = 0$. This shows that the terms in question do integrate to zero, and the different volume integrals are equal.

This is interesting because the volume expression for the charge in (73) has a simple physical meaning. Perturbatively, Q_ψ is the integral of the matter current $T_{ab}n^b$ in the direction of the CKV ξ^a . Consider a comoving cosmological observer, *i.e.* an observer whose four-velocity is the geodesic $n^a = (\partial/\partial t)^a$. The perturbative mass density measured by a comoving observer is $\delta\rho = T_{ab}n^an^b$, and $\delta M = \int_V \delta\rho_{co}$ is the comoving mass perturbation. To linear order in perturbation theory we then have

$$\delta Q_\xi = a(t)\delta M_{co}. \quad (76)$$

The explicit factor of the scale factor comes from the CKV and the choice of a comoving observer.

4.3 Convergence of Q_ψ and Q_ξ for CKV's

We have now established the equivalence of the two expressions for conserved charges associated with deSitter CKV's, Q_ψ and Q_ξ . However, we have not yet studied the finiteness properties of these charges given the boundary conditions on asymptotically deSitter spacetimes established in section (3.2). It is simplest to work in terms of Hamiltonian variables. In section (3.2), we showed that appropriate falloff conditions on the metric and extrinsic curvature on a spatial slice Σ are $q_{ij} - \tilde{q}_{ij} = \mathcal{O}(1/r)$ and $l^{ij} = \mathcal{O}(1/r^3)$, where $K^{ij} = Hq^{ij} + l^{ij}$. The $\mathcal{O}(1/r)$ corrections to the extrinsic curvature enter only from the Hq^{ij} term. Since the boundary integral expression for Q_ξ is the same for KV's and CKV's we can treat both at once. Let ξ^a be a (C)KV and decompose ξ^a as $\xi^a = Fn^a + \beta^a$, where $\beta^an_a = 0$. The canonical momentum on

the slice Σ is given by $\pi^{ij} = \sqrt{q}(K^{ij} - q^{ij}K)$. Let $p^{ij} = \pi^{ij} - \tilde{\pi}^{ij}$ be the perturbation to the canonical momentum and $\rho_{ij} = q_{ij} - \tilde{q}_{ij}$ be the perturbation to the spatial metric. The boundary term Q_ξ as defined in equation (63) can then be rewritten in terms of these Hamiltonian variables as [25]

$$Q_\xi = -\frac{1}{16\pi} \int da_i \left\{ \left(F(\tilde{D}^i \rho - \tilde{D}_j \rho^{ij}) - \rho \tilde{D}^i F + \rho^{ij} \tilde{D}_j F \right) + \frac{1}{\sqrt{\rho}} \left(\beta^i \tilde{\pi}^{jk} \rho_{jk} - 2\beta^j \tilde{\pi}^{ik} \rho_{jk} - 2\beta^j p^i_j \right) \right\} \quad (77)$$

where \tilde{D} is the covariant derivative operator on the spatial slice. For Q_ξ to be finite, we see that F must generically be independent of r , and that $|\beta|$ fall off like $1/r$, as r goes to infinity. From equation (12) we see that the only CKV for which Q_ξ will be generically finite is the generator of conformal time translations ζ^a given in equation (13). Further, we see that in general none of the charges generated by the KV's are finite. For example, the static time translation KV is given by $\chi^a = (\partial/\partial t)^a - Hx^j(\partial/\partial x^j)^a$ in cosmological coordinates, which makes the momentum term in equation (77) diverge. An exception to this is the Schwarzschild-deSitter spacetime itself. In this case, substituting the spacetime (37) into (77) leads to a cancelation between the momentum terms in equation (77). Therefore the terms which depend on β don't contribute. The charge Q_χ then turns out to be finite and given simply by $Q_\chi = M$.

4.4 Dynamical Interpretation of Q_ξ for CKV's

In the preceding sections, we have explored properties of the conserved charge Q_ζ associated with the conformal time translation CKV of deSitter and found that in certain respects it is a close parallel of the ADM mass for asymptotically flat spacetimes. In this section, we will briefly discuss another parallel, that Q_ζ is the value of the gravitational Hamiltonian if one evolves using a vector field that asymptotically approaches ζ at infinity.

In reference [15] Regge and Teitelboim studied the Hamiltonian evolution of asymptotically flat spacetimes. They showed that in order for the variational principle to be well defined, a boundary term must be added to the gravitational Hamiltonian. On solutions to the equations of motion, the volume contribution to the Hamiltonian vanishes, and the value of the Hamiltonian is given by the boundary term. If the Hamiltonian evolution is carried out along a vector field that asymptotes to the time translation Killing vector of flat spacetime, then this boundary term is simply the ADM mass.

It is straightforward to show that a similar situation holds for asymptotically deSitter spacetimes using the techniques of reference [24]. The gravitational Hamiltonian includes the volume term $H_V = \int_V \sqrt{\gamma}(F\mathcal{H} + \beta^a \mathcal{H}_a)$, where \mathcal{H} is given by the left hand side of equation (39), $\mathcal{H}_a = -2D_b \pi_a^b / \sqrt{q}$, and F, β^a are Lagrange multipliers. Let the vector $\xi^a = F n^a + \beta^a$ where n^a is the unit normal to V and $n^a \beta_a = 0$.

As discussed in reference [15], the quantity H_V is not precisely the correct functional to generate the Hamiltonian flow along the vector field ξ . A boundary term B_ξ must be added to the volume term to ensure that when one carries out the variation to derive Hamilton's equations, the total boundary term which arises in the variation actually vanishes. It was shown in [24] that finding the correct boundary term B_ξ for a general background spacetime amounts to computing the adjoint of the differential operator $(FH + \beta^a H_a)$ and keeping track of the total derivative terms. The expression found for B_ξ in [24] agrees with that given for Q_ξ in equation (77). The equivalence of the expression in Hamiltonian variables and the covariant expression in equation (63) was demonstrated in reference [25]. In the asymptotically deSitter case, choosing ξ^a to approach the conformal time translation killing vector ζ^a in the asymptotic region gives the Hamiltonian flow in conformal time. On solutions to the equations of motion H_V vanishes, and the value of the Hamiltonian is given by the value of the boundary term B_ζ , which is the same as our conserved charge Q_ζ .

One might wonder about the feasibility of defining a mass with respect to time translation in the static time coordinate, which is a symmetry of deSitter. There turn out to be several interesting drawbacks with this approach. The static time KV is not timelike everywhere and one therefore loses the positivity of the stress-energy term. To get around this one can try setting the stress-energy to zero or constraining it to be nonzero only within a horizon volume, where the KV is timelike. In Section (3) we showed that the vector fields $-\bar{\psi}\gamma^5\gamma\psi$, where ψ is a deSitter Killing spinor, are KV's rather than CKV's. In Appendix A, we give the derivation of a spinor identity in which the vector field which enters is this KV. One finds that the spinor contribution to the volume term has positive definite and negative definite contributions, both of which are in general nonzero. Of course it still could be true that the sum is always positive, but the spinor construction does not make any indications of this.

There is another potentially interesting possibility which we have not explored. Since the static KV is timelike within a horizon volume, one could study evaluating the boundary term on the deSitter horizon. This would require that the spacetime approach deSitter near the horizon in some suitable sense. Although the spinor term in the KV construction is not sign definite, as discussed in the previous paragraph, it is possible that this term is positive within a horizon volume. Abbott and Deser [6] showed that this is true to quadratic order. Further analysis would be needed to show that the gravitational contribution is positive to all orders. This choice of boundary is, of course, observer dependent. However, given the recent interest in the entropy of deSitter spacetimes defined with respect to the part of the spacetime accessible to an observer (see *e.g.* reference [26]), investigations restricted to this portion of the spacetime might prove useful.

A Spinor Identity for deSitter KV's and anti-deSitter CKV's

Generalize the Nester form of section (3) to

$$B^{ab} = \bar{\psi}(\gamma^{\hat{5}})^p \gamma^{abc} \hat{\nabla}_c \psi, \quad (78)$$

where p can be 0 or 1. One can then prove the more general identity

$$\begin{aligned} \nabla_a (B^{ab} + B^{\dagger ab}) n_b &= (G_{ab} + \Lambda g_{ab}) \xi^a n^b \\ &+ 2(-1)^p q^{ij} \hat{\nabla}_i \psi^\dagger (\gamma^{\hat{5}})^p \hat{\nabla}_j \psi - 2(-1)^p (\gamma^i \hat{\nabla}_i \psi)^\dagger (\gamma^{\hat{5}})^p (\gamma^j \hat{\nabla}_j \psi) \\ &+ (-1)^p H(-c + (-1)^p c^*) \left[\psi^\dagger (\gamma^{\hat{5}})^p \gamma^j \hat{\nabla}_j \psi - \right. \\ &\left. (-1)^p (\gamma^j \hat{\nabla}_j \psi)^\dagger (\gamma^{\hat{5}})^p \psi \right], \end{aligned} \quad (79)$$

where $\xi^a = -\bar{\psi}(\gamma^{\hat{5}})^p \gamma^a \psi$ and n^b is the unit normal vector to a spacelike slice. When ψ is a Killing spinor, then as discussed in section (2) for $p = 0(1)$, ξ^a a deSitter CKV (KV) or an anti-deSitter KV (CKV). For $p = 0$, equation (79) reduces to the result of section (3). For $p = 1$ the volume term, which is quadratic in $\hat{\nabla} \psi$, is easily seen to be no longer sign definite. Decompose ψ into eigenvectors of $\gamma^{\hat{5}}$,

$$\psi = \frac{1}{2}(1 - \gamma^{\hat{5}})\psi + \frac{1}{2}(1 + \gamma^{\hat{5}})\psi \equiv \psi_- + \psi_+. \quad (80)$$

The second term on the right hand side is then given by

$$2q^{ij}(-\nabla_i \psi_+^\dagger \nabla_j \psi_+ + \nabla_i \psi_-^\dagger \nabla_j \psi_-). \quad (81)$$

The Dirac-Witten equation written in terms of the projected spinors is given by $\gamma^j \nabla_j \psi_\mp + i\frac{H}{2}(n-1)\psi_\pm = 0$, which has no solutions with $\psi_+ = 0$ and nontrivial ψ_- . From this we see that the volume integrand will include negative, as well as positive contributions. It is still possible that the overall result could be overall positive as argued in reference [16], but the spinor construction does not appear to give any indication of this.

We now want to check that the charges Q_ψ and Q_ξ agree for the deSitter Killing vectors, as was the case for the charges corresponding to the deSitter CKV's. To do this, we start again by dividing the Nester form as in equation (32) and follow the analysis of the boundary term as in section (3.2) above. The terms which are linear in H , which canceled for $p = 0$, now add for $p = 1$. There is then an additional contribution to the boundary term in the KV case. Using the perturbative result $\delta\gamma_a = \frac{1}{2}h_{ab}\tilde{\gamma}^b$, the boundary term becomes

$$Q_\psi = Q_\psi|_{p=0} + i\frac{H}{16\pi} \int ds_{ab} \bar{\tilde{\psi}} \gamma^{\hat{5}} (h\tilde{\gamma}^{ab} - \tilde{\gamma}^{ac} h_c^b + \tilde{\gamma}^{bc} h_c^a) \tilde{\psi}, \quad (82)$$

where $Q_\psi|_{p=0}$ is given in (50). The $p = 0$ term only depends on the spinor fields through the vector contribution $\xi^a = -\bar{\psi}\tilde{\gamma}^a\tilde{\psi}$. The additional boundary term in the

$p = 1$ case depends on the product of two gamma matrices sandwiched between $\tilde{\psi}$ and $\tilde{\psi}$. As one can see from equation (21), this quantity is proportional to the shear $w_{ab} = \nabla_{[a}\xi_{b]}$ of the Killing field ξ^a . The shear contribution of this shear term to Q_ψ then matches the contribution to Q_ξ coming from the shear term in equation (61).

References

- [1] M. S. Turner, “Making sense of the new cosmology,” [arXiv:astro-ph/0202008].
- [2] E. Witten, “A Simple Proof Of The Positive Energy Theorem,” Commun. Math. Phys. **80**, 381 (1981).
- [3] P. Breitenlohner and D. Z. Freedman, “Positive Energy In Anti-De Sitter Backgrounds And Gauged Extended Supergravity,” Phys. Lett. B **115**, 197 (1982).
- [4] P. Breitenlohner and D. Z. Freedman, “Stability In Gauged Extended Supergravity,” Annals Phys. **144**, 249 (1982).
- [5] G. W. Gibbons, C. M. Hull and N. P. Warner, “The Stability Of Gauged Supergravity,” Nucl. Phys. B **218**, 173 (1983).
- [6] L. F. Abbott and S. Deser, “Stability Of Gravity With A Cosmological Constant,” Nucl. Phys. B **195**, 76 (1982).
- [7] G. W. Gibbons and C. M. Hull, “A Bogomolny Bound For General Relativity And Solitons In N=2 Supergravity,” Phys. Lett. B **109**, 190 (1982).
- [8] S.D. Majumdar, Phys. Rev. **72**, 930 (1947).
- [9] A. Papapetrou, Proc. Roy. Irish Acad. **A51**, 191 (1947).
- [10] E. Witten and D. I. Olive, “Supersymmetry Algebras That Include Topological Charges,” Phys. Lett. B **78**, 97 (1978).
- [11] D. Kastor and J. Traschen, “Cosmological multi - black hole solutions,” Phys. Rev. D **47**, 5370 (1993) [arXiv:hep-th/9212035].
- [12] K. Pilch, P. van Nieuwenhuizen and M. F. Sohnius, “De Sitter Superalgebras And Supergravity,” Commun. Math. Phys. **98**, 105 (1985).
- [13] K. i. Nakao, T. Shiromizu and K. i. Maeda, “Gravitational mass in asymptotically de Sitter space-times,” Class. Quant. Grav. **11**, 2059 (1994).
- [14] R. M. Wald, “Black Hole Entropy In Noether Charge,” Phys. Rev. D **48**, 3427 (1993) [arXiv:gr-qc/9307038].
- [15] T. Regge and C. Teitelboim, “Role Of Surface Integrals In The Hamiltonian Formulation Of General Relativity,” Annals Phys. **88**, 286 (1974).

- [16] T. Shiromizu, D. Ida, and T. Torii , “Gravitational energy, dS/CFT correspondence, and cosmic no-hair,” [arXiv:hep-th/0109057].
- [17] A. Strominger, “The ds/CFT correspondence,” JHEP **0110**, 034 (2001) [arXiv:hep-th/0106113].
- [18] D. Kastor and J. Traschen, “Particle production and positive energy theorems for charged black holes in De Sitter,” Class. Quant. Grav. **13**, 2753 (1996) [arXiv:gr-qc/9311025].
- [19] L. J. Romans, “Supersymmetric, cold and lukewarm black holes in cosmological Einstein-Maxwell theory,” Nucl. Phys. B **383**, 395 (1992) [arXiv:hep-th/9203018].
- [20] J. Nester, “A New Gravitational Expression with a Simple Positivity Proof,” Phys. Lett. **83A**, 241 (1981).
- [21] R. M. Wald, “General Relativity,” *University of Chicago Press* (1984).
- [22] J. Lee and R. D. Sorkin, “Derivation Of A Bogomol’ny Inequality In Five-Dimensional Kaluza-Klein Theory,” Commun. Math. Phys. **116**, 353 (1988).
- [23] G. McVittie, Mon. Not. R. Astron. Soc. **93** (1933) 325.
- [24] J. Traschen, “Constraints On Stress Energy Perturbations In General Relativity,” Phys. Rev. D **31**, 283 (1985).
- [25] L. F. Abbott, J. Traschen and R. Xu, “Integral Constraints In General Relativity,” Nucl. Phys. B **296**, 710 (1988).
- [26] R. Bousso, “Adventures in de Sitter space,” [arXiv:hep-th/0205177].