

2004

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Recommended Citation

Yun, Min; Reddy, NA; Scoville, NZ; Frayer, DT; Robson, EI; and Tilanus, RPJ, "Multiwavelength Observations of the Gas-rich Host Galaxy of PDS 456: A New Challenge for the ULIRG-to-QSO Transition Scenario" (2004). *The Astrophysical Journal*. 1130. 10.1086/380559

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MULTI-WAVELENGTH OBSERVATIONS OF THE GAS-RICH HOST GALAXY OF PDS 456: A NEW CHALLENGE FOR THE ULIRG-TO-QSO TRANSITION SCENARIO

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to appear in the Astrophysical Journal

ABSTRACT

We report new K -band, radio continuum, and CO (1–0) imaging observations and 850 μm photometric observations of PDS 456, the most luminous QSO in the local universe ($z < 0.3$). The 0.6'' resolution K -band image obtained using the Keck telescope shows three compact $m_K \sim 16.5$ ($M_K \sim -21$) sources at a projected distance of ~ 10 kpc to the southwest, and the host galaxy of PDS 456 may be interacting or merging with one or more companions. The observations using the OVRO millimeter array has revealed a narrow CO (1–0) line (FWHM = 181 km s⁻¹) centered at $z = 0.1849$, and $9 \times 10^9 M_\odot$ of molecular gas mass is inferred. Radio continuum luminosity is nearly an order of magnitude larger than expected from its FIR luminosity, and the radio source, unresolved by the 2'' beam of the VLA, is dominated by the AGN activity. Our 850 μm photometric observations suggest that the cold dust content of the host galaxy is less than one half of the amount in Arp 220.

The analysis of the spectral energy distribution reveals both a QSO-like and a ULIRG-like nature, and the observed IR, X-ray, and gas properties suggest that the AGN activity dominates its luminosity. PDS 456 displays many characteristics expected of an object undergoing a transition from an ultra-luminous infrared galaxy (ULIRG) to a classical QSO phase as proposed by Sanders et al., including an optical spectrum dominated by broad emission lines, large X-ray and IR luminosity, a large cold gas/dust content, and an extremely large $L_{\text{FIR}}/M_{\text{H}_2}$ ratio ($\gtrsim 100 L_\odot/M_\odot$). ULIRGs and IR QSOs form a broad continuous track in the “star formation efficiency” plot in the manner consistent with the ULIRG-QSO transition scenario, relating the evolution in the dust processed luminosity with the available fuel (gas and dust) supply. However, the location of PDS 456 is clearly offset from the apparent track traced by the ULIRGs and IR QSOs on this plot. Therefore, PDS 456 appears to be a rare, exceptional object, and the duration of the physical process governing its present properties must be short compared with the length of the luminous QSO phase.

Subject headings: galaxies: active — galaxies: starburst — infrared: galaxies — radio lines: galaxies — quasars: individual (PDS 456, IRAS 08572+3915)

1. INTRODUCTION

The remarkable similarity between cosmic star formation history (e.g. Madau et al. 1996) and quasar evolution (e.g. Shaver et al. 1996) suggests an intriguing possibility of coeval evolution between galaxies and massive black holes in the universe. This is strengthened by the observed correlation between black hole mass and mass of host galaxy’s spheroid component (Magorrian et al. 1998; Ferrarese & Merritt 2000; Gebhardt et al. 2000). The bulk of star formation and AGN activity in the early universe may have been driven by mergers of massive, dusty, and gas-rich galaxies (see Blain et al. 1999), and the high frequency of submm continuum detections of high redshift optically selected QSOs (about 30%, see Omont et al. 2001; Carilli et al. 2001) also support the coeval formation and evolution

scenario.

If the same underlying physical process continues to drive the formation and evolution of massive galaxies and black holes today, then evidence for such activities may still be observable. Citing both starburst and AGN activities along with large molecular gas contents among the most luminous infrared galaxies in the local universe, Sanders et al. (1988a) suggested an evolutionary connection between ultraluminous infrared galaxies (ULIRGs) and QSOs. Possible formation scenarios for supermassive black holes inside ULIRGs have been suggested by several theoretical investigations (Weedman 1983; Norman & Scoville 1988; Taniguchi, Ikeuchi, & Shioya 1999). Ubiquitous presence of one or more luminous active galactic nucleus among ULIRGs has been questioned recently (see Gen-

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zel et al. 1998; Tacconi et al. 2002), but nevertheless a close link between the QSO and the ULIRG phenomenon appears to exist and deserves further investigations.

Here, we report a detailed multi-wavelength observational study of PDS 456 (IRAS 17254–1413), the most luminous QSO in the local universe ($z \lesssim 0.3$), in order to shed further light on the ULIRG-QSO connection. As part of a systematic survey for young stellar objects at Pico dos Dias Observatory, Torres et al. (1997) have obtained an optical spectrum of the 15th magnitude stellar object PDS 456, selected based on its bright IRAS detection ($S_{25\mu\text{m}} = 750$ mJy, $S_{60\mu\text{m}} = 930$ mJy). To their surprise, its optical spectrum is dominated by broad emission lines characteristic of a QSO at a redshift of $z = 0.184$. When corrected for extinction, its absolute B magnitude of -26.7 is 30% more luminous than 3C 273, the most luminous quasar in the local universe previously known. Because it is located near the Galactic Plane ($b = +11^\circ$), confusion is a serious problem at $100 \mu\text{m}$, and only an upper limit of $S_{100\mu\text{m}} \leq 1$ Jy is offered by the IRAS data. Nevertheless, both $S_{25\mu\text{m}}/S_{60\mu\text{m}} = 0.81$ and $S_{60\mu\text{m}}/S_{100\mu\text{m}} \geq 0.93$ suggest PDS 456 is a warm IRAS source, consistent with the presence of a luminous AGN (see de Grijp et al. 1985; Yun et al. 2001). Its $60 \mu\text{m}$ luminosity alone is $\nu L_\nu \sim 10^{12} L_\odot$, high even among the ultraluminous infrared galaxies in the local universe. Therefore presence of significant amount of molecular gas and associated massive starburst is a strong possibility although not all warm ultraluminous IRAS sources are detected in CO emission (see below).

Given its large bolometric and IR luminosity, PDS 456 represents a particularly important object, potentially representing a critical transition between an ultraluminous infrared galaxy and a *bona fide* QSO with large optical/UV luminosity and broad emission lines. We have investigated the nature of the host galaxy by obtaining a high resolution image using the Keck telescope in the *K*-band where the QSO-host contrast may be the most favorable for revealing the underlying galaxy. We have also imaged PDS 456 using the VLA in order to locate the non-thermal activity with subarcsecond accuracy and CO (1–0) emission using the OVRO millimeter array in order to determine whether the host galaxy contains a large amount of molecular gas, typical of ULIRGs. Finally, submillimeter continuum measurements have also been obtained using the SCUBA camera on James Clerk Maxwell Telescope (JCMT) in order to constrain the thermal dust spectrum and dust mass. We examine these new observational results in terms of the ULIRG-QSO evolution scenario.

2. OBSERVATIONS

2.1. *K*-band Imaging at the Keck telescope

Near infrared observations of PDS 456 were obtained on September 1, 2001 with the Near Infrared Camera (NIRC) on the W. M. Keck telescope (Matthews & Soifer 1994). Observations were carried out using the *K* filter centered at $2.21 \mu\text{m}$ with a FWHM of $0.43 \mu\text{m}$. The final image was constructed from the 8 dithered images with 30 second exposure time each, for a total exposure of 240 seconds. Seeing was $0.6''$ as measured from several point sources in the field. The dithered images were dark-subtracted, flat-fielded, and then combined us-

ing the NIRCtools package written by D. Thompson and run using the Image Reduction and Analysis Facility (IRAF).² The fractional offsets used to combine the images were determined from the IRAF task IMALIGN which centroids on several point sources in the field common to all of the images. The resulting rotated *K*-band image is shown in Figure 1.

2.2. 8.5 GHz and 1.2 GHz Imaging at the VLA

Radio continuum image of PDS 456 at 8.5 GHz (3.6 cm) wavelength was obtained using the Very Large Array (VLA)³ in Socorro, NM, on January 13, 1998. The observations were made in the D-configuration using 14 antennas in the continuum mode with 100 MHz total bandwidth. The nearby quasar 1730–130 was observed periodically to track the instrumental gain, and 3C 286 was used for absolute flux calibration. The data were calibrated and imaged following the standard procedure within the NRAO AIPS software system. The naturally weighted image has an rms noise of $1\sigma = 21 \mu\text{Jy}$. PDS 456 is detected with a peak flux density of 4.5 ± 0.9 mJy ($P_{8.5\text{GHz}} = (3.5 \pm 0.7) \times 10^{23} \text{ W Hz}^{-1}$)⁴, unresolved by the synthesized beam of $13'' \times 8''$ (PA= -29°) in size.

Radio continuum image of PDS 456 at 1.2 GHz (25 cm) wavelength were obtained using the VLA on July 16, 1999. The observations were made in the A-configuration using 27 antennas in the spectral line mode with 6.25 MHz total bandwidth. The nearby quasar 1730–130 was observed periodically to track the instrumental gain, and 3C 286 was used for absolute flux calibration. The naturally weighted image has an rms noise of $1\sigma = 0.26$ mJy. PDS 456 is detected with a peak flux density of 24 ± 5 mJy ($P_{1.2\text{GHz}} = (1.9 \pm 0.4) \times 10^{24} \text{ W Hz}^{-1}$), unresolved by the synthesized beam of $4'' \times 2''$ (PA= 22°) in size. The best astrometric position of PDS 456 is derived from this image, which is $\alpha(J2000) = 17^{\text{h}} 28^{\text{m}} 19^{\text{s}}.78$ and $\delta(J2000) = -14^\circ 15' 56''.12$ with an uncertainty of less than $0''.4$.

2.3. CO (1–0) Imaging at OVRO

Aperture synthesis CO (1–0) observations of PDS 456 were carried out using the Owens Valley Millimeter Array between April 2000 and May 2001. There are six 10.4 m diameter telescopes in the array, providing a field of view of about $75''$ (FWHM) at 100 GHz. The telescopes are equipped with SIS receivers cooled to 4 K, and typical single sideband system temperatures (T_{SSB}) were between 250 and 350 K. Baselines of 15–115 m E-W and 15–115 m N-S were used. The total on-source integration time was about 28 hours.

The redshifted CO (1–0) line at $z = 0.184$ occurs at 97.3574 GHz. A digital correlator configured with 120×4 MHz channels (12 km s^{-1}) covered a total velocity range of 1440 km s^{-1} . By tuning the receivers to this frequency, the spectrometer covers the redshift range of $z = 0.1811 - 0.1869$. The nearby quasar 1730–130 was observed at 20 minute intervals to track the phase and short term instrument gain. Uranus ($T_B = 120$ K), Neptune ($T_B = 115$ K), 3C 273, and 3C 454.3 were observed for the absolute flux calibration. The data were calibrated using the standard Owens Valley array program MMA (Scoville et al. 1992) and were mapped and analyzed using the imaging program DIFMAP (Shepherd et al. 1994) and the NRAO AIPS

² IRAF is written and supported by the IRAF programming group at the National Optical Astronomy Observatories (NOAO) in Tucson, Arizona. NOAO is operated by the Association of Universities for Research in Astronomy (AURA), Inc. under cooperative agreement with the National Science Foundation.

³ The National Radio Astronomy Observatory is a facility of the National Science Foundation operated under cooperative agreement by Associated Universities, Inc.

⁴ The quoted uncertainty in the flux density reflects the 20% uncertainty in the transfer of the absolute flux scaling.

software system. The uncertainty in absolute flux calibration is about 20%, mainly due to the uncertainty in transferring the calibration between the source and the flux calibrators. The positional accuracy of the resulting maps is better than $\sim 0''.5$.

2.4. 850 μm Photometry

The 850 μm photometry observations were conducted on October 26 & 29, 2002, using the submillimeter camera SCUBA on the James Clerk Maxwell Telescope (JCMT) in a photometry mode. The 225 GHz zenith opacity was 0.1 and 0.06, respectively, on those days, and high temporal resolution measurements from the JCMT line-of-sight water-vapor radiometer were used to correct for the opacity variations. The flux calibration was done using Uranus. Due to the early times in the night, the atmosphere was somewhat unstable. This was compounded by the low elevation ($EL \sim 30^\circ$) of the source at the time of the observations. The sensitivity achieved is about 2.5 mJy (1σ) after a total of 2 hours of integration. The data does not reveal any positive emission above the noise level, but the 3σ upper limit of 7.5 mJy at 850 μm provides a useful constraint on the dust continuum spectrum and the amount of cold dust associated with the host galaxy of PDS 456.

3. RESULTS

3.1. *K*-band Image

The final co-added *K*-band image of PDS 456 is shown in Figure 1. The point-like QSO and the point spread function (PSF) of the telescope dominates the image, and three compact sources lying about $3''$ to the southwest of the QSO are seen (named K1, K2, and K3). We performed aperture photometry on the QSO as well as the three compact components, and the results are summarized in Table 1. Photometric zero-points are determined using separate observations of several standard stars observed at varying airmass throughout the night. Aperture corrections were made for the 5, 7, and 20 pixel radius aperture for K1, K2+K3, and the QSO, respectively. We derive *K*-band magnitude of 9.89 ± 0.02 for the QSO. After subtracting the wings of the QSO, we derive a *K*-band magnitude of 16.3 ± 0.2 for K1. K2 and K3 are approximately equal in brightness, and the combined *K*-band magnitude is 15.6 ± 0.1 . The near-IR photometry for the QSO reported by the Two Micron All Sky Survey (2MASS: Skrutskie et al. 1997) is $m_J = 12.03 \pm 0.02$, $m_H = 11.08 \pm 0.02$, and $m_K = 9.86 \pm 0.02$ (see Table 2), in an excellent agreement with our measurement. The absolute astrometry for the *K*-band image is good to only about $1''$, and the astrometric information is improved by replacing the the *K*-band QSO centroid position with the radio centroid position (§ 2.2).

The whole co-added *K*-band image is about $50''$ by $50''$ in size and includes about 90 sources brighter than $K = 20$. The integrated source count is shown in Figure 2. The number density of sources brighter than $K = 16.5$ is about 10 per square arcminute, and this corresponds to about 3×10^4 objects per square degree, which is nearly two orders of magnitudes higher than the *K*-band galaxy count in the field (e.g., see Fig. 2 by Djorgovski et al. 1995). A natural explanation for the elevated source count is that PDS 456 is located in the direction of Galactic Center, near the Galactic Plane, at $l = 10^\circ.4$ and $b = +11^\circ.2$. Therefore foreground stars likely account for most of the sources found in the *K*-band image.

The three compact features K1, K2, and K3 found adjacent to the QSO are very likely physically associated with PDS 456,

and the host galaxy may have undergone a tidal interaction or a merger recently. For $H_0 = 75 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.3$, and $\Omega_\lambda = 0.7$, angular distance for PDS 456 at $z = 0.185$ is 595 Mpc, and the angular separation of $3''$ corresponds to a projected distance of 9 kpc, which is the natural size scale for the underlying host galaxy or companions. The source count shown in Figure 2 suggests that the probability of a random $K = 16.5$ source falling within a $3''$ radius of any given position is less than 1%. The likelihood of a chance projection for *three* random sources of this magnitude within $3''$ of PDS 456 is less than 10^{-6} . Therefore, a random occurrence of three foreground or background sources near the position of PDS 456 is extremely unlikely. If they are at the same distance as PDS 456, their *K*-band absolute magnitude (M_K) ranges between -20.7 and -21.1 , which is about 4 magnitudes fainter than that of an L^* galaxy. Attempts to subtract the QSO image using the PSF derived from calibration stars were not successful in revealing any underlying structures within ~ 4 kpc of the QSO ($M_K = -27.5$). An L^* stellar host may be present underneath the QSO image, but it cannot be determined by our present data.

A faint extension to the southwest of the QSO is also visible even in the Digitized Sky Survey and 2MASS plate images. This feature is not clearly resolved by these low resolution images, however, and little else can be learned from them. In an unpublished *J*-band image obtained with the Pico dos Dias Observatory 1.6 m telescope in 1999 by F. Jablonski under a good seeing condition, sources K1 and K2+K3 appear clearly resolved, and the K1 source, which appears extended toward the QSO, has a *J*-band magnitude of 17.8 ± 0.2 while the combined brightness of K2+K3 is 17.5 ± 0.1 (F. Jablonski & C. Torres 2003, private communication). The *J*–*K* colors of K1 and K2+K3, 1.5 ± 0.3 and 1.9 ± 0.2 respectively, are somewhat redder than the ensemble average color of 2MASS galaxies, $\langle J-K \rangle = 1.2$, but they are consistent with that of a stellar system located at the redshift of the QSO when the foreground extinction (see § 4.2) and the *k*-correction are taken into account. In an imaging survey of 20 nearby luminous QSOs at $z < 0.3$ using the HST, Bahcall et al. (1997) reported finding eight companion galaxies with *V*-band magnitude within 4 magnitudes of L^* at projected distances less than 10 kpc. Other HST imaging studies have also reported a high frequency of tidal tails and close companions associated with the nearby luminous QSOs (McLeod & McLeod 2001; Dunlop et al. 2003). Therefore finding compact stellar systems surrounding PDS 456 itself is not unusual, but the large number and the asymmetric distribution are curious. The asymmetry and the clumpy appearances are similar to the features seen around 3C 273 (Martel et al. 2003) and PKS 2349–014 (Bahcall et al. 1997; McLure et al. 1999), and therefore these stellar clumps are also likely the remnants of a recent accretion or a merger event. A definitive characterization of the nature of these clumps will require a more sensitive higher angular resolution imaging using the HST.

3.2. CO (1–0) and 100 GHz Continuum

The CO (1–0) channel maps of PDS 456 are shown in Figure 3. The naturally weighted image has a synthesized beam of about $7.0'' \times 4.8''$ (PA= -9°). Emission features brighter than 6 mJy beam^{-1} (4σ) are seen within $3''$ of the QSO position marked with a cross at channels corresponding to 97.269 GHz and 97.301 GHz. The CO (1–0) spectrum obtained at the QSO position after smoothing to the 16 MHz (49 km s^{-1}) and 64 MHz (197 km s^{-1}) resolution is shown in Figure 4. The thick

solid curve is a model Gaussian line profile with $\sigma = 25$ MHz (77 km s^{-1} ; $\text{FWHM} = 181 \text{ km s}^{-1}$) centered at 97.283 GHz ($z = 0.1849$). The velocity integrated line flux is $1.45 \pm 0.29 \text{ Jy km s}^{-1}$. At a luminosity distance of 834 Mpc, this translates to a H_2 mass of $9 \times 10^9 M_\odot$ using the standard conversion factor derived from the Galactic molecular clouds (see Young & Scoville 1991). This derived H_2 mass is on the high end of the mass range found for PG QSOs ($1\text{--}10 \times 10^9 M_\odot$, Evans et al. 2001; Scoville et al. 2003) and is only slightly less than those of ULIRGs ($2\text{--}5 \times 10^{10} M_\odot$ using the same conversion, e.g. Solomon et al. 1997).

The spatial resolution of the CO data is about $21 \text{ kpc} \times 15 \text{ kpc}$, and the CO emitting structure is not expected to be substantially resolved. However, the centroids of the brightest CO peaks in the two channel maps shown in Figure 3 are offset from each other by about $2''$ (6 kpc). Given the low S/N of the data, this displacement of the peak positions may not be very significant. On the other hand, the displacement of the blueshifted CO peak is in the direction of the two K -band sources K2 & K3 (see Fig. 1). These two K -band sources are likely associated with the host of the QSO (see § 3.1), and this apparent alignment with the CO displacement, perhaps fortuitous, adds to the possibility that the molecular gas associated with PDS 456 may be extended over 5-10 kpc scales.

The data from the image sideband of the OVRO receivers are also cross-correlated to produce continuum visibility data centered at 100.3574 GHz. The 100 GHz continuum image obtained from these visibility data has an rms noise of $0.2 \text{ mJy beam}^{-1}$, and no statistically significant source is found with the field of view of the interferometer. From this we can place an upper limit of 0.6 mJy (3σ) for the QSO continuum flux density at 100 GHz. The continuum spectral energy distribution of PDS 456 is discussed in greater detail below.

4. DISCUSSIONS

4.1. Spectral Energy Distribution of PDS 456

The spectral energy distribution (SED) holds information on the temperature and the physical processes involved, even when the angular resolution of the observations are insufficient to yield any structural information. In examining the ULIRG-QSO connection, a comparison of the SEDs is particularly useful since the underlying physical processes and host galaxy properties can be directly compared.

The SED of PDS 456, shown with large squares in Figure 5, nicely demonstrates its QSO-like and ULIRG-like nature. The SEDs of the previously most luminous QSO 3C 273 ($z = 0.158$), an IR luminous radio-quiet QSO I Zw 1 ($z = 0.061$), and the prototypical ULIRG Arp 220 ($z = 0.018$) are shown after redshifting their continuum spectra to $z = 0.185$ for a direct comparison. The SED data for PDS 456 are also tabulated in Table 2. The SED data for 3C 273, I Zw 1, and Arp 220 come from the NASA Extragalactic Database (NED) and literature. The SEDs of PDS 456 and 3C 273 are remarkably similar and are essentially identical between the far-IR (FIR) and optical wavelengths. When corrected for extinction of $A_V \sim 1.5$, the bolometric luminosity of PDS 456 becomes about 30% larger than that of 3C 273 (Torres et al. 1997; Simpson et al. 1999). Identified as one of the “warm ULIRGs” that are possibly in transition from galaxy to quasar by Sanders et al. (1988b), 3C 273 has both thermal and non-thermal emission mechanisms contributing to its SED (see Courvoisier 1998). The SEDs at radio wavelengths differ by more than 3 orders of magnitude, reflect-

ing the presence of a beamed radio jet in 3C 273. This part of the SED contributes little to the total luminosity, however.

The SED of PDS 456 covering the far-IR through radio wavelengths is modeled using the dusty starburst SED model by Yun & Carilli (2002) and is shown using a solid curve in Figure 5. As in QSOs 3C 273 and I Zw 1, non-thermal emission from the AGN dominates the short wavelength ($\lambda < 50 \mu\text{m}$) part of the SED in PDS 456, but dust reprocessed AGN emission and light from young stars should account for the dust peak from millimeter to FIR wavelengths (e.g. Rowan-Robinson 1995). There are no continuum detections of PDS 456, only upper limits, between 100 GHz ($3000 \mu\text{m}$) and 2000 GHz ($150 \mu\text{m}$), and therefore there is only limited information to constrain the cold gas content and cold dust properties. The turnover in the FIR dust peak occurs at a much higher frequency, near 6000 GHz ($50 \mu\text{m}$), implying the dust temperature is on average much higher than that in Arp 220 ($T_d \sim 45 \text{ K}$, Scoville et al. 1991; Yun & Carilli 2002). The model SED shown in Figure 5 is chosen to match the $100 \mu\text{m}$ measurement by Reeves et al. (2000) and has dust temperature of 120 K in its rest frame. This means a substantial amount of dust in the host galaxy of PDS 456 is exposed to a significantly higher mean radiation field and to the high energy photons originating from the AGN. The total dust mass derived from the existing SED data, assuming $T_d = 120 \text{ K}$ and emissivity $\beta = 1.5$, is $0.4 \times 10^7 M_\odot$ (using Eq. 3 of Yun & Scoville 1998). Both the dust mass and temperature are well within the range of dust properties derived for a large sample of Palomar-Green (PG) QSOs by Haas et al. (2000, $T_d = 20\text{--}120 \text{ K}$, $M_d = 10^{7\pm 1} M_\odot$). The effective diameter for the emitting area is about 135 pc (using Eq. 4 of Yun & Carilli 2002), which is similar in size to the nuclear starburst region in Arp 220 (Scoville, Yun, & Bryant 1997; Sakamoto et al. 1999). The strong mid-IR emission from PDS 456 and 3C 273 are indicative of abundant warm and hot dust (a few hundred to 1000 K), and this is a clear indication of an energetic AGN in these objects (see Figure 10 of Yun et al. 2001). The measured radio continuum flux density of PDS 456 is 8 times larger than what is expected from a starburst system with the same FIR luminosity, and this is a clear indication that a radio AGN is also present in PDS 456.

The total dust mass of PDS 456 could be potentially much larger if substantial amount of cold ($T_d \sim 10\text{--}40 \text{ K}$) dust is present. The best constraint on the cold dust content come from our new $850 \mu\text{m}$ measurement made using the SCUBA camera on JCMT. As shown in Figure 5, the dust continuum emission from Arp 220 should have been detected by our observation with $\text{S/N} \gtrsim 5$. The 3σ upper limit at $850 \mu\text{m}$ shown here suggests that the total cold dust mass of PDS 456 is less than one half of the amount in Arp 220 ($T_d \sim 45 \text{ K}$). The model SED for PDS 456 shown in Figure 5 is consistent with that of the composite starburst+QSO system I Zw 1 with estimated cold dust mass of $(1\text{--}6) \times 10^7 M_\odot$ (Hughes et al. 1993; Haas et al. 1998; Andreani, Franceschini, & Granato 1999). More sensitive submillimeter continuum measurements are needed to constrain the total dust mass in PDS 456.

A useful insight into the AGN contribution to the observed FIR and bolometric luminosity can be obtained by examining its hard X-ray luminosity. The best-fit model spectrum for the BeppoSAX measurements yields an intrinsic 2-10 keV luminosity of about $5.3 \times 10^{44} \text{ erg s}^{-1}$ (Vignali et al. 2000, after correcting for H_0), which is about 10% of its total FIR luminosity, $L_{\text{FIR}} = 1.3 \times 10^{12} L_\odot$. If the AGN activity should account for a

substantial fraction of its FIR and bolometric luminosity, one would expect the X-ray luminosity of the AGN to be comparable or larger than the FIR luminosity. While the intrinsic 2-10 keV luminosity does not account for the entire FIR luminosity, this fraction is quite substantial, similar to other IR bright, optically selected QSOs and Seyfert 1 AGNs and 2-3 orders of magnitudes larger than most ULIRGs and composite starburst+AGN systems (see Figure 6; also Risaliti et al. 2000; Levenson, Weaver, & Heckman 2001). PDS 456 stands out even among the FIR bright QSOs as its $L_{2-10\text{keV}}/L_{\text{FIR}}$ ratio is an order of magnitude larger than that of IR luminous QSOs Mrk 1014 and I Zw 1 and two orders of magnitudes larger than that of a prototypical ULIRG/QSO system Mrk 231. Therefore, the hard X-ray properties of PDS 456 is much closer to optically selected luminous QSOs, and it can be strongly argued that the bolometric luminosity of PDS 456 is dominated by the AGN activity.

4.2. Geometry of the Circum-AGN Disk

Most of the warm dust responsible for the mid- and far-IR emission modeled in Figure 5 has to be located within a rotationally supported circum-AGN disk surrounding the central AGN in order to achieve the observed high mean temperature. A lower limit to the dynamical mass of the gas/dust disk surrounding the optical QSO in PDS 456 can be estimated using the FIR source size derived in § 4.1 and the gas rotation speed derived from the CO line width in Figure 4. Assuming the gas is in a circular rotation with rotation speed of $V_c = 90 (\sin i)^{-1}$ km s⁻¹ at a radius of 70 pc, the resulting dynamical mass is

$$M_{\text{dyn}} = 1.3 \times 10^8 \left[\frac{R}{70 \text{ pc}} \right] \left[\frac{V_c}{90 \text{ km/s}} \right]^2 (\sin i)^{-2} M_{\odot}.$$

This is far smaller than the gas mass inferred from the CO luminosity, $9 \times 10^9 M_{\odot}$ (see § 3.2), unless the gas disk is viewed nearly face-on with an inclination angle smaller than $i \lesssim 8^{\circ}$. The CO-to-H₂ conversion factor is thought to be 2-3 times smaller than the standard Galactic value for the warm, dense clouds in circum-nuclear regions (Scoville, Yun, & Bryant 1997; Downes & Solomon 1998) while both V_c and R are only known to a factor of two or so. However, these factors alone cannot fully account for the difference between the inferred molecular gas mass and the dynamical mass.

The simple fact that the optical QSO is readily visible requires the line-of-sight to the QSO be mostly free of any obscuring material, and this also favors a face-on geometry for the circum-AGN disk. A similar inference on the face-on nature and the resolution of the apparent discrepancy between the derived gas mass and the dynamical mass for gas disks within ULIRGs hosting an optical QSO (e.g. Mrk 231) has been made previously (Bryant & Scoville 1996; Downes & Solomon 1998). Since PDS 456 is seen through the Galactic plane, some foreground extinction is naturally expected, and the spectral slope between the near-IR and optical wavelengths is indeed steeper than expected. The analysis of the optical and near-IR colors, continuum spectral shape, and the Balmer line ratios all consistently suggests reddening of the QSO light by $A_V \sim 1.5$ (see Torres et al. 1997; Simpson et al. 1999). Torres et al. further estimate that the reddening due to the Galactic foreground derived from the extinction map of Burstein & Heiles (1982) and from the diffuse interstellar bands and Na D1 line in their spectra also suggest a foreground reddening of $A_V \sim 1.5$, accounting for *all* of the reddening associated with the QSO light. The only result that contradicts this conclusion is the

analysis of the X-ray spectrum that requires a highly ionized absorber with $N_{\text{H}}(\text{warm}) \sim 5 \times 10^{24}$ cm⁻² and an additional cold absorber with $N_{\text{H}}(\text{cold}) \sim 3 \times 10^{22}$ cm⁻² (Vignali et al. 2000). However, this inference is highly model dependent.

The CO emission from PDS 456 may instead arise primarily from the cold gas and dust yet undetected in continuum, rather than from the circum-AGN disk traced in IR in Figure 5. The channel maps shown in Figure 3 suggests that the CO emitting region may be extended over a region 10 kpc in extent (see § 3.2). Even with the physical extent of 5-10 kpc for the CO emitting region, the dynamical mass M_{dyn} is still smaller than the H₂ mass inferred from the CO luminosity unless the inclination of the disk is more face-on ($i \lesssim 30^{\circ}$).

While the face-on geometry for the dust and gas disks in PDS 456 and Mrk 231 seems quite secure, a surprising result is that a general inference of a face-on geometry for nuclear gas disks in other QSO host systems is *not* supported by the observed distribution CO line widths. As demonstrated by the discussion above, the face-on geometry requirement arises mainly from the small observed line widths in the dynamical mass calculation. However, when the histograms of observed CO line widths are compared between QSO hosts and ULIRGs as shown in Figure 7, QSO hosts show comparable CO line widths as ULIRGs. The median CO line width for ULIRGs is larger ($\Delta V \sim 300$ km s⁻¹) than that of QSO hosts ($\Delta V \sim 250$ km s⁻¹), consistent with the general expectation, but the difference is much smaller than expected from the viewing geometry consideration alone. Furthermore, the CO line widths observed in three out of nine QSO hosts are larger than 350 km s⁻¹ (FWHM), too large to be consistent with a face-on geometry. One solution to this general problem is that the parsec scale circum-AGN disk that dictates the viewing angle of the central source is not aligned with the extended gas/dust disk or ring traced in CO. Since only a few of these objects have been imaged in CO or in dust continuum with sub-kpc resolution, the possibility of optical AGNs viewed through gaps in a large scale gas/dust ring or disk at an arbitrary inclination angle or through a highly warped disk cannot be ruled out. Therefore, Figure 7 cautions strongly against the common assumption of a face-on geometry for gas and dust distribution around optically visible QSOs.,

4.3. Dusty QSO hosts and ULIRGs

In the ULIRG-to-QSO evolutionary scenario proposed by Sanders et al. (1988a), an ultraluminous infrared galaxy is thought to go through a warm IR source phase as observed in PDS 456 today with large IR luminosity, warm IR color, a large molecular gas content, and a dust enshrouded luminous AGN. In accordance with this scenario, a large number of optically selected PG quasars that are also detected by IRAS have been studied in detail and are found to contain large amounts of molecular gas (Alloin et al. 1992; Evans et al. 2001; Scoville et al. 2003). Similarly, the majority of warm ULIRGs are also shown to be composite objects hosting one or more luminous AGNs and an active starburst fueled by the large molecular gas content in their nuclear regions (Sanders et al. 1988b).

IR luminosity and total gas/dust mass are the two key variables characterizing the ULIRG-QSO evolution. Specifically, the transition from a ULIRG to a QSO is driven by rapid conversion of gas into stars and/or the subsequent growth of a massive black hole, followed by the dispersion of gas and dust surrounding these activities. Therefore such an evolution scenario

should follow a distinct track in the plot between IR luminosity and the total gas mass. In particular, a dusty QSO emerging from a ULIRG phase should show the characteristic large FIR luminosity and rapidly diminishing gas content, distinguishable from ULIRGs. If the IR-bright QSO PDS 456 is such a transition object, then it should appear in the bridging region between the areas occupied by ULIRGs and by QSOs.

To scrutinize the ULIRG-QSO evolutionary scenario further, optically selected IR QSOs, including PDS 456, are compared with ULIRGs and luminous IR galaxies (LIRGs) in a plot of FIR luminosity (L_{FIR}) versus CO luminosity (L'_{CO}), which can be translated into a total molecular gas mass (M_{H_2}) using the standard conversion relation (see Young & Scoville 1991). This plot shown in Figure 8 is commonly referred to as the “star formation efficiency (SFE)” plot because historically it is used to demonstrate that IR-bright starburst systems are not only forming stars at high rates but also with a greater efficiency, producing 1-2 orders of magnitudes more luminosity (thus more massive stars) per solar mass of molecular gas. A constant ratio of $L_{FIR}/M_{H_2} = 100 L_{\odot}/M_{\odot}$ is an upper bound to what is expected if massive star formation is primarily responsible for the FIR luminosity (see Scoville & Soifer 1991), and this plot can also offer a glimpse on whether the dust obscured powering source is massive stars or has to be a dust enshrouded AGN. Among the 14 optically selected QSOs with measured molecular gas contents (shown in filled circles), only Mrk 1014 is found in the area of high FIR luminosity and a large molecular gas mass occupied by ULIRGs (empty circles). The remaining 13 QSOs fall near the $L_{FIR}/M_{H_2} = 10 L_{\odot}/M_{\odot}$ line which is characteristic of less luminous IR starbursts (crosses; Sanders et al. 1991). This comparison suggests that *optically selected QSOs are distinct from the ULIRG population in general, even when their host galaxies are fairly rich in molecular gas* ($M_{H_2} = 10^{9-10.5} M_{\odot}$).

The location of PDS 456 is distinct from most ULIRGs, less luminous IR starbursts, and even other optically selected QSOs in Figure 8. The uncertainty associated with the data points plotted in Figure 8 is typically about 20%, and the size of the error bars should be comparable or smaller than the size of the symbols used. Therefore the observed scatter in this plot reflects substantial variations in the underlying physical processes (see below), far in excess of the measurement uncertainties. PDS 456 has about 5 times more FIR luminosity than other IR detected PG QSOs and LIRGs with comparable CO luminosity (molecular gas mass). In fact, it is one of the objects showing the most extreme L_{FIR}/M_{H_2} (L_{FIR}/L'_{CO}) ratio, and such a high L_{FIR}/M_{H_2} ratio is reasonably expected if a large fraction of the FIR luminosity arises from the ionizing radiation originating from the AGN. Another noteworthy object with a similarly extreme L_{FIR}/M_{H_2} ratio is IRAS 08572+3915, which is a well-known warm ULIRG ($S_{25\mu m}/S_{60\mu m} = 0.23$ and $S_{60\mu m}/S_{100\mu m} = 1.64$) with a Seyfert/LINER optical spectrum and shows IR-excess (see Sanders et al. 1988b; Veilleux et al. 1995; Yun et al. 2001). From the near-IR and mid-IR imaging and spectroscopy, Dudley & Wynn-Williams (1997), Imanishi & Dudley (2000), and Soifer et al. (2000) have argued that dust-obscured AGN activity is the dominant energy source for IRAS 08572+3915. On the other hand, the failure to detect hard X-ray emission ($L_{2-10keV} < 4.4 \times 10^{41}$ erg s⁻¹, Risaliti et al. 2000) or an unresolved VLBI radio source (Smith et al. 1998) challenges the AGN interpretation. Both X-ray and radio signatures of an AGN can be obscured by a large column

of neutral and ionized gas, and we identify its proximity to the $L_{FIR}/M_{H_2} = 100 L_{\odot}/M_{\odot}$ line in Figure 8, adjacent to the *bona fide* QSO PDS 456, as another evidence that AGN activity is the primary power source for IRAS 08572+3915.

The rare and unique nature of PDS 456 is further demonstrated by the fact that the majority of the “warm ULIRGs”, identified as possible ULIRGs in transition to a QSO phase by Sanders et al. (1988b), are indistinguishable from other ULIRGs in Figure 8. Nine out of the 12 “warm ULIRGs” identified by Sanders et al. have been observed in CO thus far. All nine have been detected in CO, and they are already included as ULIRGs in Figure 8. All but one of these sources cluster together with the other ULIRGs in this figure. IRAS 08572+3915 is the one exception, and indeed it stands out from the others with a large L_{FIR}/M_{H_2} ratio as noted already. No CO measurements are available for the remaining three (IRAS 01003–2238, IRAS 12071–0444, 3C 273), and we cannot rule out the possibility that one or more of these three objects also have the large L_{FIR}/M_{H_2} ratio similar to PDS 456. Regardless, we can safely conclude that objects like PDS 456 are rare, even among the “warm ULIRGs”.

Morphological studies of QSO host galaxies have found that not all QSO hosts display evidence of a recent merger while IR-detected QSO hosts often do (see Bahcall et al. 1997; Clements 2000; Surace, Sanders, & Evans 2001). Because the size of the accretion disk, mass accretion rate, and total gas mass requirement for QSO activity are relatively small, major mergers involving two massive gas-rich galaxies is not a necessary ingredient for the QSO phenomenon in general. However, given the high probability of finding more massive black holes in more massive progenitors (Magorrian et al. 1998; Ferrarese & Merritt 2000; Gebhardt et al. 2000) involved in major mergers as well as the theoretical possibility of forming and/or growing massive black holes within massive merger remnants (Weedman 1983; Norman & Scoville 1988; Taniguchi, Ikeuchi, & Shioya 1999), the likelihood of finding one or more massive black holes in an object undergoing a ULIRG phase should be naturally quite high. To this end, a recognizable trend is expected in Figure 8 between ULIRGs and IR-detected QSOs. Indeed ULIRGs (empty circles) and IR QSOs (filled circles) form a continuous distribution in Figure 8 – IR QSOs follow a diagonal $L_{FIR}/M_{H_2} \sim 20 L_{\odot}/M_{\odot}$ line, which connects smoothly to the nearly vertical distribution of ULIRGs near $\log M_{H_2} \sim 10.5$. If an IR-detected QSO represents a later stage of evolution following the ULIRG phase, then this trend would indicate that the post-ULIRG evolution occurs along a constant L_{FIR}/M_{H_2} line characteristic of IR bright starbursts.

This observed trend between ULIRGs and IR QSOs is somewhat surprising since it is not what is predicted by the existing ULIRG-QSO evolution scenarios. There is a plausible *a posteriori* explanation to this trend, however. Some scatter is seen in the distribution of the ULIRGs, LIRGs, and PG QSOs in Figure 8, and one can interpret this as evidence for a wide range of “efficiency” in converting fuel into luminosity. Alternatively, one can interpret that this wide range of observed “efficiency” reflects a range in the ratios of dust processed luminosity contribution by AGN and starburst activity in these objects. Among the optically visible QSOs, as radiation pressure and winds clear out gas and dust from the immediate surroundings and reveal the optical AGN, the geometrical cross-section and the solid angle for dust heating by the AGN decrease rapidly (see Downes & Solomon 1998), and the AGN contribution to

the FIR luminosity diminishes quickly to the level where the underlying massive star formation again dominates the FIR luminosity. In this scenario, a nearly vertical upturn in the distribution of ULIRGs near $L_{CO} \sim 10^{10} \text{ K km s}^{-1} \text{ pc}^2$ represents ubiquitous presence of luminous AGNs among ULIRGs, providing a broad range of additional luminosity. Arp 220, which is thought to be powered mostly by an intense starburst, thus appears in the lower half of the ULIRG distribution while Mrk 231 and Mrk 1014, both hosting luminous AGNs, appear near the top of the vertical ULIRG distribution where the elevated AGN contribution to the bolometric luminosity becomes significant. Relatively little overlap between the two populations suggests that the transition is quite rapid.

As noted earlier, PDS 456 and IRAS 08572+3915 are clearly displaced from ULIRGs and IR QSOs in Figure 8, and they represent an interesting challenge to the above ULIRG-QSO evolutionary scenario. One possible explanation is that these objects follow a slightly different evolutionary path. Starting initially as a ULIRG with one or more luminous AGNs (e.g. Mrk 231, Mrk 1014), one possible evolutionary path they might follow is a rapid exhaustion of the gas reservoir via a yet unknown mechanism, retaining only the compact, dense, and warm circum-AGN dust cocoon – i.e., evolving nearly horizontally to the left in Figure 8. This evolutionary scenario is qualitatively similar to what was proposed by Sanders et al. (1988a) although little detailed descriptions were offered by these authors. Few objects are found along this possible evolutionary track, however, and this is not likely a commonly followed evolutionary path.

An alternative explanation is that objects like PDS 456 represent a brief transient phase of highly elevated FIR emission in the ULIRG-QSO evolution. One possible scenario is a momentary obscuration of the AGN by a episodic inflow of gas and dust, resulting in brief periods of dust coverage with a large solid angle and brief increases in FIR luminosity by a factor of a few to ten. Such an event would materialize in Figure 8 as a vertical upward displacement anywhere along the nominal evolutionary track. The rarity of objects like PDS 456 suggests that this period is brief compared with the duration of the IR QSO phase. Numerical simulations of major mergers involving two gas-rich disk galaxies (e.g. Barnes & Hernquist 1996; Mihos & Hernquist 1996) have shown that a large fraction of gas can be concentrated into the central 100 pc of the post-merger potential, fueling luminous starbursts seen in the centers of ULIRGs. These simulations also show that the gas inflow, which can last over several hundred million years after the merger, can be lumpy, predicting a significant variation in the gas arrival rate over the entire duration of merger that can last up to a billion years. If a luminous AGN is already present, then this episodic nature of the gas (and dust) inflow would predict a highly variable obscuration of the central AGN – the crossing time for a large gas/dust clump should be much shorter ($< 10^{6-7}$ years) than the galaxy merger time scale. In this scenario, objects like PDS 456 and IRAS 08572+3915 are thus seen at a special moment in their ULIRG-QSO evolution, rather than being a critical transition phase as considered in § 1. Another important consequence of this scenario is that it predicts a rare but distinct class of FIR sources with $L_{FIR}/M_{H_2} \sim 100 L_{\odot}/M_{\odot}$, spanning perhaps the entire range of L_{FIR} and M_{H_2} . A small number of “IR-excess” objects (characterized by a high FIR-to-radio flux density ratio) identified by Yun et al. (2001) among the IRAS detected galaxies may indeed represent this rare population. As mentioned already, IRAS 08572+3915 is such an object. An-

other of the few well studied “IR-excess” objects, NGC 4418, has a high SFE of $L_{FIR}/M_{H_2} = 69 L_{\odot}/M_{\odot}$ ($\log M_{H_2} = 9.17$, $\log L_{FIR} = 11.01$, Sanders et al. 1991) and is indeed thought to host a dust-enshrouded AGN (see Spoon et al. 2001).

The unusually high L_{FIR}/M_{H_2} ratio observed for objects like PDS 456 and IRAS 08572+3915 may have a different explanation, unrelated to the ULIRG-QSO evolution scenario. Whatever the process, it has to account for both the large IR luminosity and the large dust/gas content simultaneously. The paucity of objects with the extreme L_{FIR}/M_{H_2} ratio of PDS 456 suggests that it also has to be a rare process, and its duration must be quite short compared with the length of the luminous phase for QSOs.

5. SUMMARY

We report a new, detailed multi-wavelength observational study of PDS 456, the most luminous QSO in the local universe ($z \lesssim 0.3$). The 0.6'' resolution *K*-band image obtained using the Keck telescope shows three compact $m_K \sim 16.5$ ($M_K \sim -21$) sources located ~ 10 kpc southwest of the QSO, but little detail on the host galaxy is revealed. The near-IR color of these objects are consistent with their being at the same redshift as PDS 456. These imaging results suggest that the host galaxy of PDS 456 may be interacting or merging with one or more companions.

Observations using the OVRO millimeter array has revealed a narrow CO (1–0) line ($\sigma = 77 \text{ km s}^{-1}$; FWHM = 181 km s^{-1}) centered at $z = 0.1849$. The narrow line width is consistent with a face-on geometry for the circum-AGN gas disk as also suggested by the optically visible QSO. A total H_2 mass of $9 \times 10^9 M_{\odot}$ is inferred for the host galaxy from the CO luminosity. The most accurate astrometry of the QSO is derived with subarcsecond accuracy using the new high resolution VLA continuum image at 1.2 GHz, and new radio continuum and millimeter continuum measurements are used to constrain the long wavelength SED of PDS 456. The eight times larger radio continuum flux density than expected from the radio-FIR correlation is a clear indication that a radio AGN is present in PDS 456.

Our 850 μm photometric observations of PDS 456 using the SCUBA camera on JCMT did not produce a detection at the level of 7.5 mJy (3σ). On the other hand, this upper limit is sufficient to rule out an Arp220-like host with copious amount of cold dust for PDS 456. The total amount of cold dust present in the host system does not exceed about one half of the dust present in Arp 220.

The SEDs of PDS 456, I Zw 1, and 3C 273 are remarkably similar between optical and FIR wavelengths, including the pronounced mid-IR enhancement which is a well known characteristic of AGNs (de Grijp et al. 1985; Yun et al. 2001). At the same time, the SED between FIR and radio wavelengths is similar to that of an ultraluminous galaxy Arp 220, suggesting a dual QSO-like and ULIRG-like nature. The observed FIR, X-ray, and CO properties of PDS 456 suggest that the AGN activity dominates its luminosity.

While a face-on geometry for the dust and gas disks in PDS 456 and Mrk 231 seems quite secure, a surprising new result is that a broad inference of a face-on geometry for nuclear gas disks among QSO host systems is *not* supported by the observed distribution CO line widths. Among the small number of ULIRGs and PG QSOs examined, QSO hosts show comparable CO line widths as ULIRGs, and the possibility of optical AGNs

viewed through gaps in a large scale gas/dust ring or disk at an arbitrary inclination angle cannot be ruled out. We conclude that a face-on geometry for gas and dust distribution around optically visible QSOs, as expected from the AGN “unification” scheme, may be an over-simplification.

PDS 456 and IRAS 08572+3915 are two objects with an extremely large L_{FIR}/M_{H_2} ratio, $\sim 100 L_{\odot}/M_{\odot}$, which is the upper bound of what can be reasonably expected from a dust enshrouded starburst. A comparison of the gas mass-to-light ratio clearly shows that optically selected IR QSOs are distinct from the ULIRG population in general, even when their host galaxies are fairly rich in molecular gas ($M_{H_2} = 10^{9.0-10.5} M_{\odot}$). Relatively little overlap between the two populations suggests that the transition is quite rapid. A continuous track formed by ULIRGs and IR QSOs in the “star formation efficiency” plot (Figure 8) suggests a transition in the luminosity contribution by dust obscured AGN from a starburst, in the context of the ULIRG-QSO evolution scenario. We further infer that the FIR luminosity of the IR QSOs is mostly dominated by the underlying starburst activity under the ULIRG-QSO evolution scenario. We also propose that objects like PDS 456 represent a special transient phase in the ULIRG-QSO evolution scenario where the central AGN is momentarily obscured by an episodic inflow of gas and dust. As a consequence, we predict a rare but distinct population of dust obscured AGNs with $L_{FIR}/M_{H_2} \gtrsim 100 L_{\odot}/M_{\odot}$ covering an entire range of FIR luminosity.

In summary, PDS 456, the most luminous QSO in the local universe, does not possess the FIR and gas properties consistent with the ULIRG-QSO evolution scenario proposed by Sanders et al. (1988b) despite its ULIRG-like FIR and CO luminosity and classical QSO signatures of broad emission lines and large

X-ray luminosity. The continuous distribution of ULIRGs and IR QSOs in the “star formation efficiency” plot (Figure 8) is broadly consistent with a ULIRG-QSO evolution scenario, and the FIR emission associated with the IR QSOs are dominated by the underlying starburst under this scenario. PDS 456 clearly deviates from the observed ULIRG-QSO trend in this plot. The paucity of PDS456-like objects in the local universe strongly suggests that PDS 456 is a rare, exceptional object, and the duration of the physical process governing its properties must be short compared with the length of the luminous QSO phase.

The authors are grateful to J. Lowenthal for his assistance with the initial examination of the Keck *K*-band imaging data. The authors also thank the Keck Observatory staff who made the *K*-band observations possible. This publication makes use of data products from the Two Micron All Sky Survey, which is a joint project of the University of Massachusetts and the Infrared Processing and Analysis Center/California Institute of Technology, funded by the National Aeronautics and Space Administration and the National Science Foundation. This research has made use of the NASA/IPAC Extragalactic Database (NED) which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration. Support for this work was provided in part by a grant from the K. T. and E. L. Norris Foundation and NSF grant AST 99-81546. N. Reddy was supported by the National Science Foundation Graduate Research Fellowship. The JCMT is operated by the Joint Astronomy Centre, on behalf of the UK Particle Physics and Astronomy Research Council, the Netherlands Organisation for Scientific Research, and the National Research Council of Canada.

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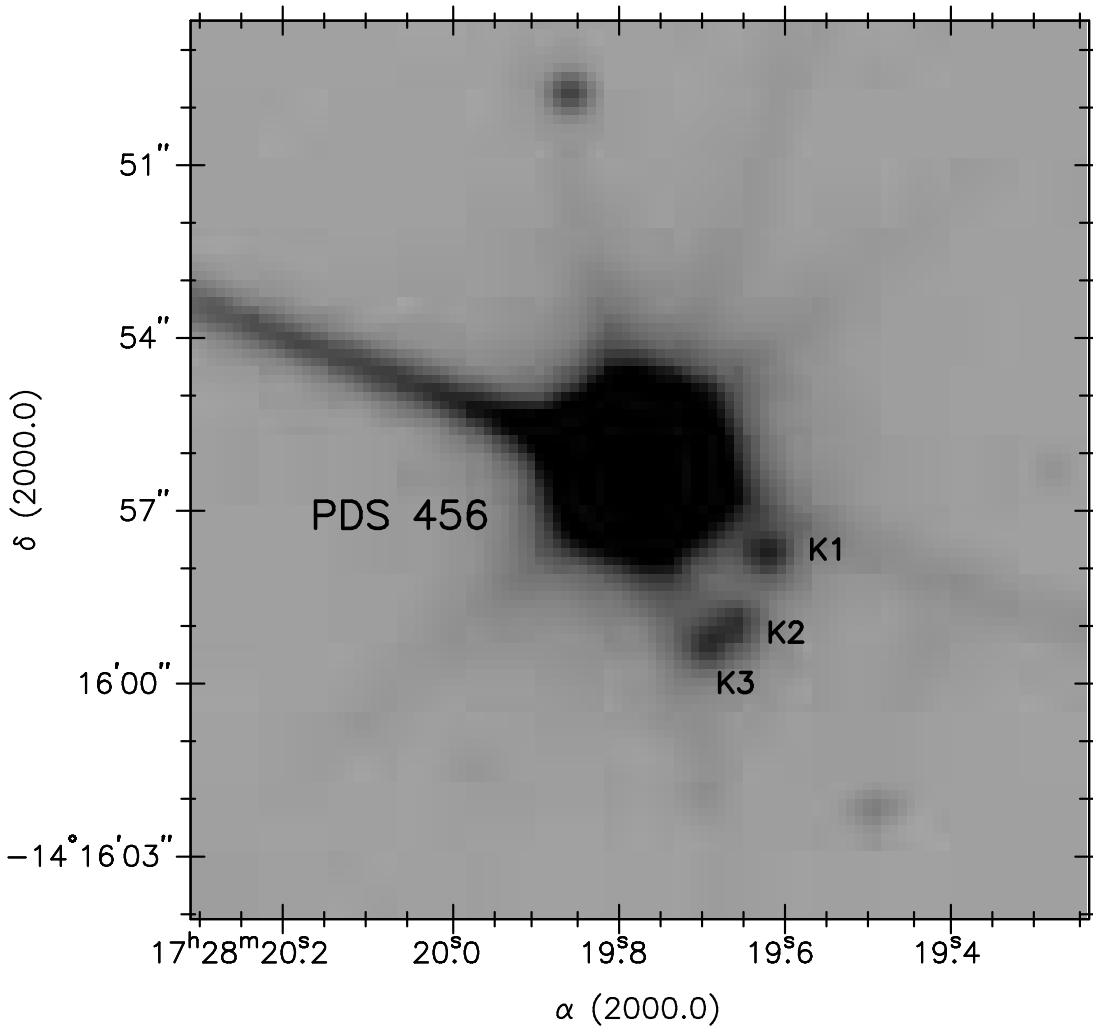


FIG. 1.— Greyscale K-band image of PDS 456 obtained at the Keck Telescope.

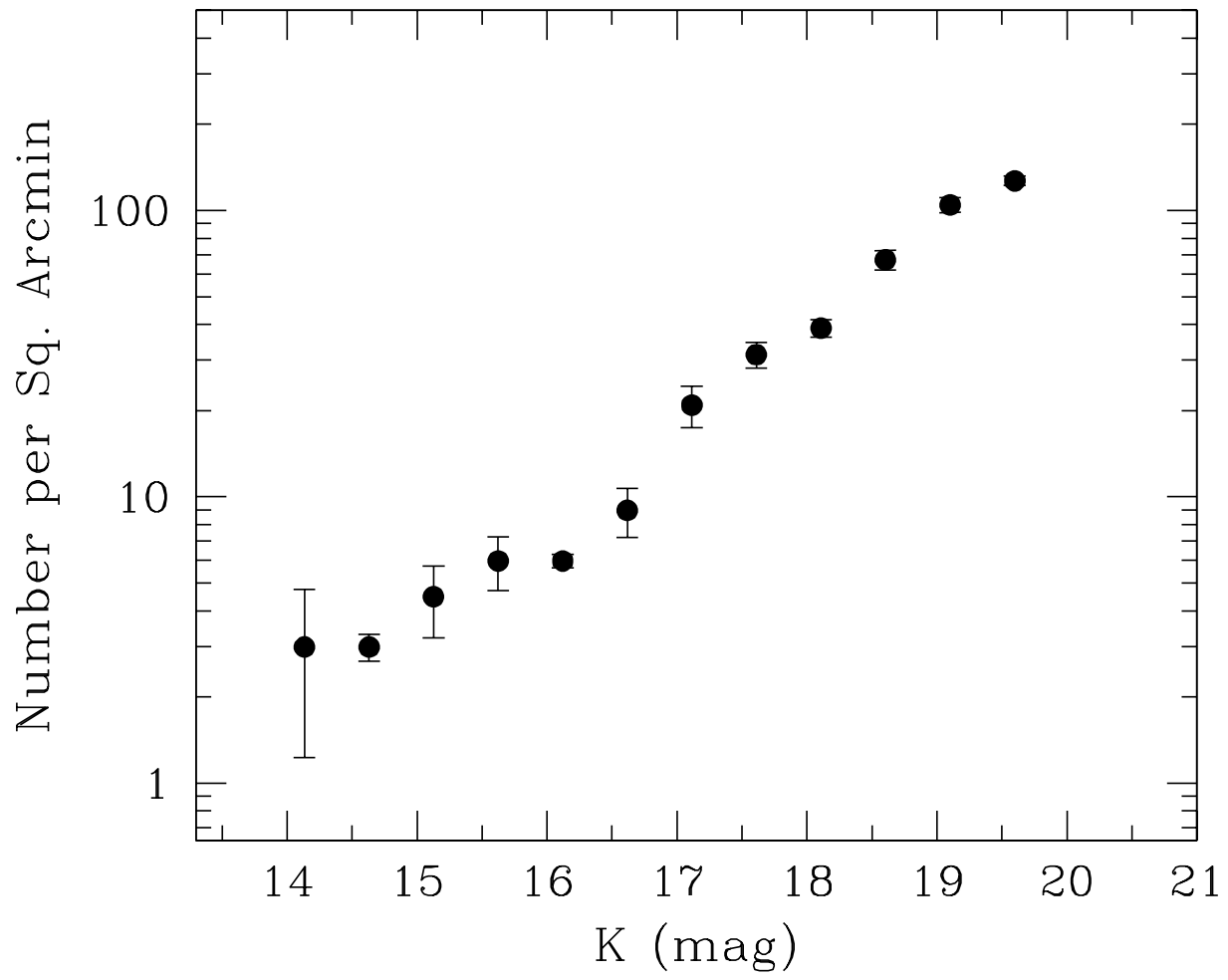


FIG. 2.— Cumulative number counts of K -band sources in the PDS 456 field. Stars dominate the brightest sources.

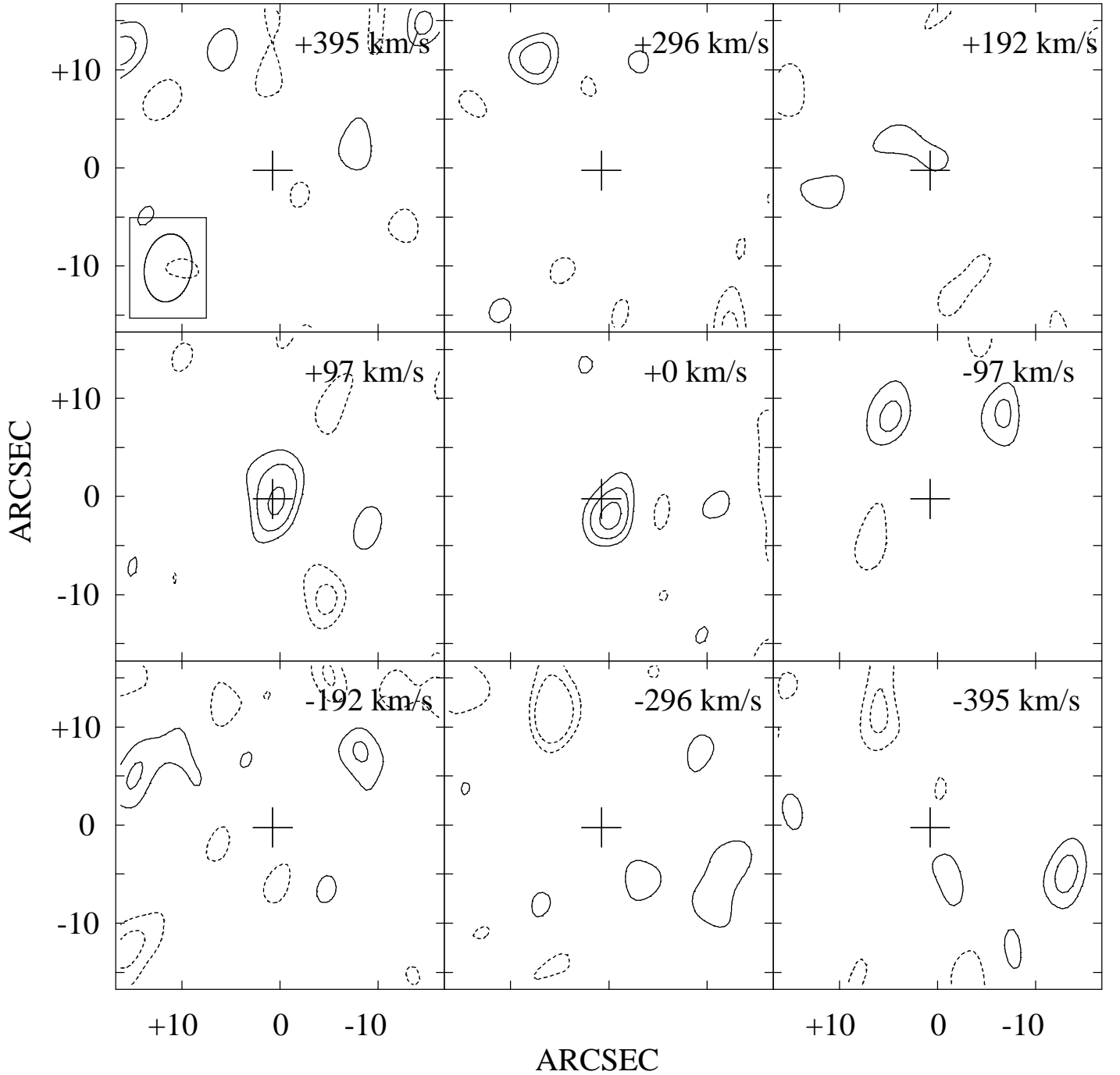


FIG. 3.— CO (1-0) channel map of PDS 456. The contour levels correspond to $-3, -2, +2, +3, +4,$ and $+5$ times $1.5 \text{ mJy beam}^{-1}$ (1σ). The velocity offsets listed are with respect to the channel with the strongest line feature at 97.301 GHz. The cross marks the optical QSO position.

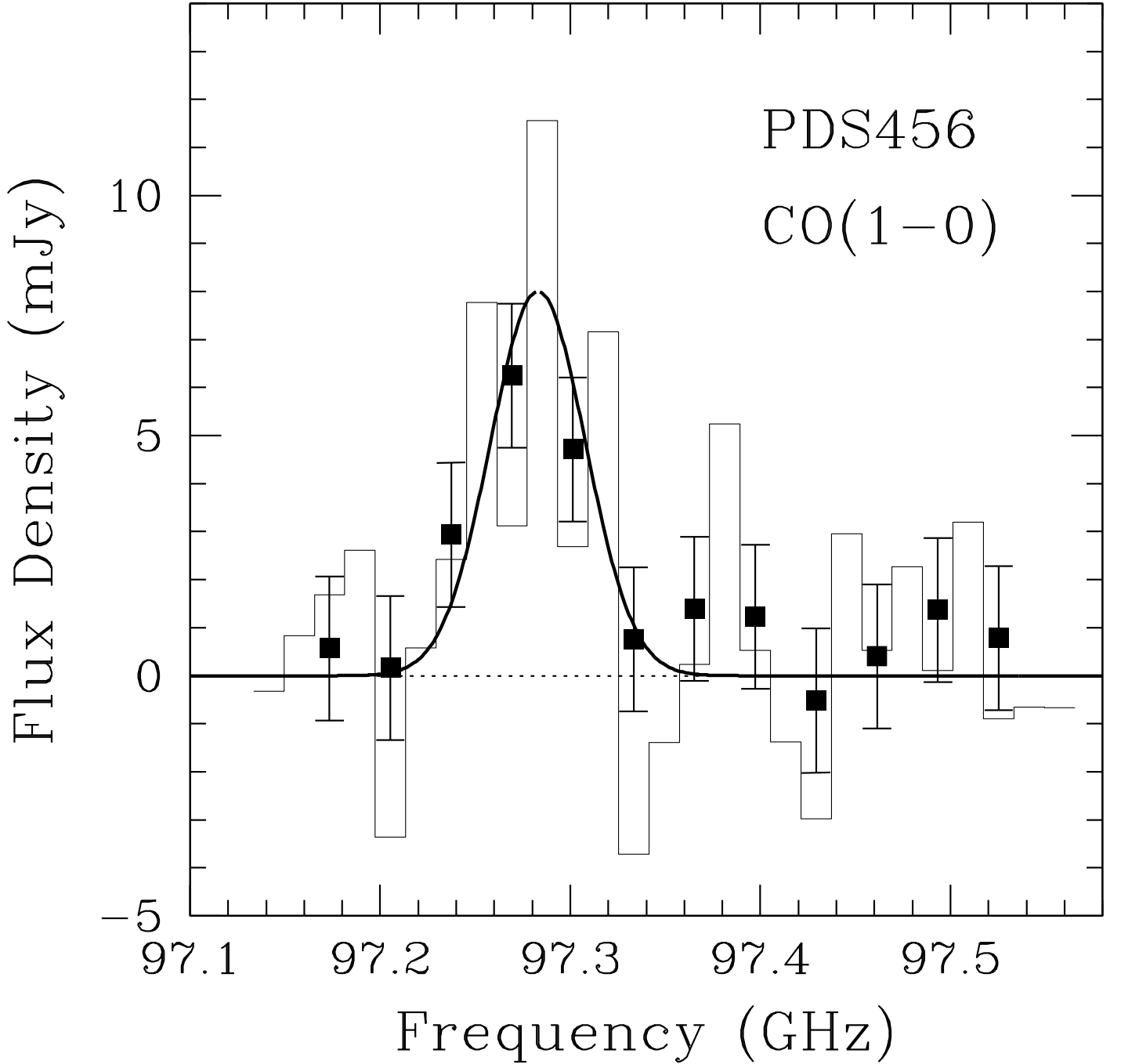


FIG. 4.— CO (1–0) spectrum of PDS 456. The histogram is the spectrum obtained at the QSO position after smoothing to the 16 MHz (49 km s^{-1}) resolution. The CO spectrum further smoothed to a 64 MHz (197 km s^{-1}) resolution is shown using filled squares with an increment of 32 MHz. The thick solid curve is a model Gaussian line profile with $\sigma = 25 \text{ MHz}$ (77 km s^{-1} ; FWHM = 181 km s^{-1}) centered at 97.283 GHz ($z = 0.1849$).

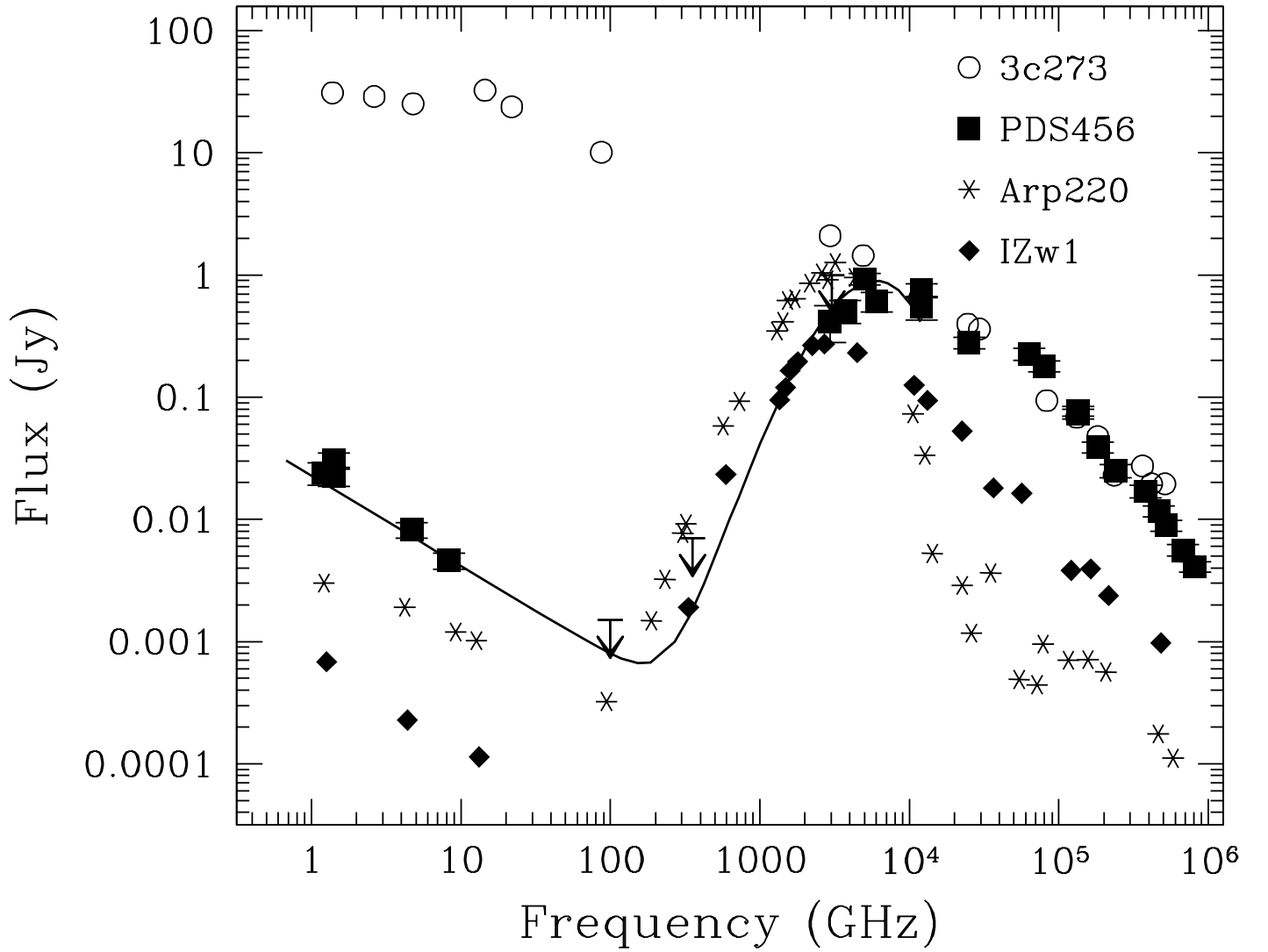


FIG. 5.— Spectral Energy Distribution (SED) of PDS 456. The SEDs of 3C 273 ($z = 0.158$), I Zw 1 ($z = 0.061$), and Arp 220 ($z = 0.018$) are also shown for comparison after redshifting to that of PDS 456 ($z = 0.185$). The solid line is a model SED for a source with dust temperature of $T_d = 120$ K, emissivity $\beta = 1.5$, and a FIR emission region with an effective diameter of 135 pc (see Yun & Carilli 2002). The radio continuum part of the model SED is a power-law with a slope of $\alpha = +0.75$, and the flux density is boosted by a factor of eight over the value predicted by the radio-FIR correlation for star forming galaxies in order to match the observed data points.

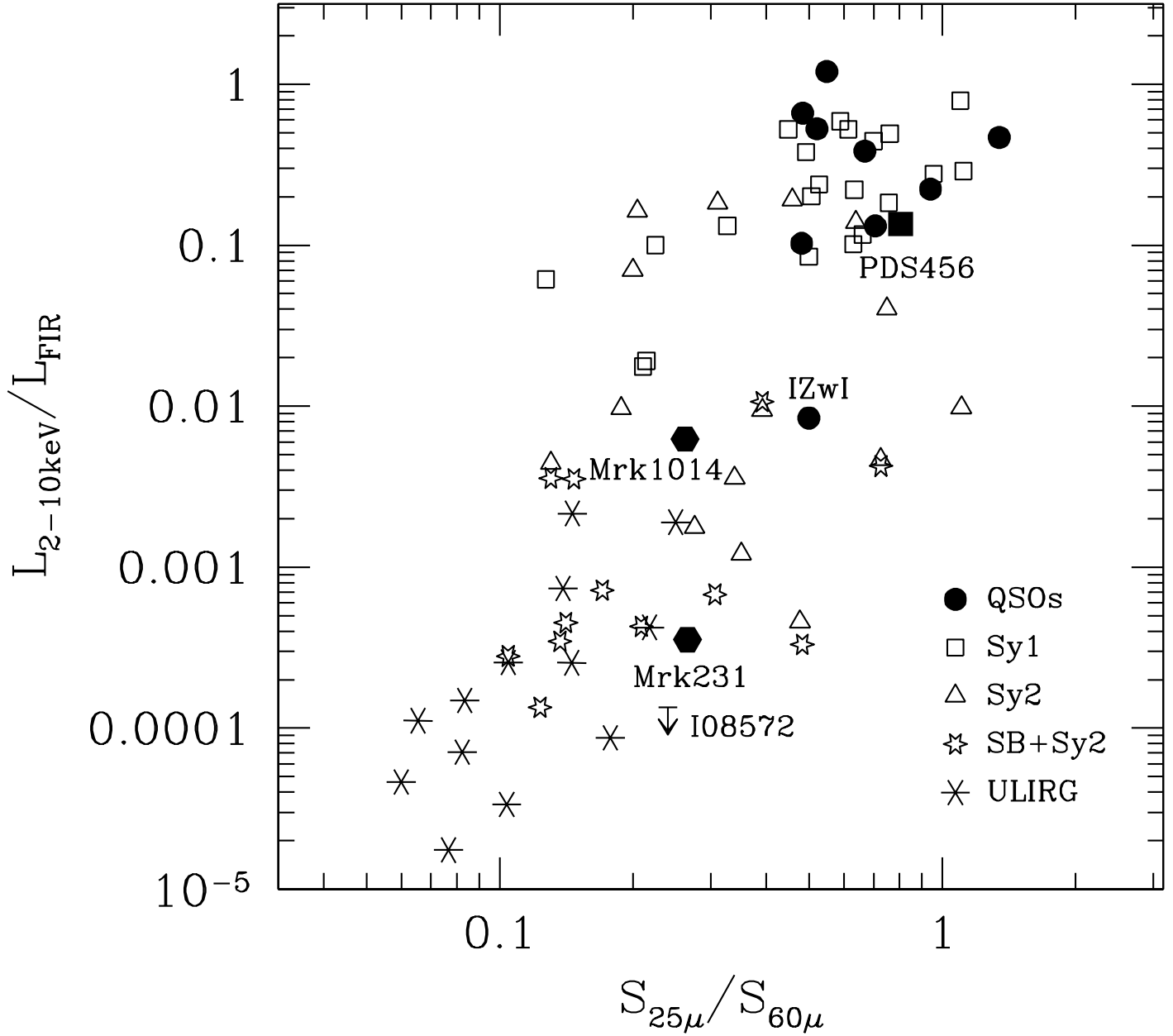


FIG. 6.— $L_{2-10\text{keV}}/L_{\text{FIR}}$ vs. $S_{25\mu}/S_{60\mu}$. The X-ray data plotted come from Mas-Hesse et al. (1995), Vignali et al. (2000), Risaliti et al. (2000), Levenson, Weaver, & Heckman (2001), and Ptak et al. (2003). Three luminous IR QSOs Mrk 231, Mrk 1014, and I Zw 1 are identified along with PDS 456 for comparison. The warm IRAS source IRAS 08572+3915 is not detected in hard X-ray ($L_{2-10\text{keV}} < 4.4 \times 10^{41}$ erg s $^{-1}$, Risaliti et al. 2000), and an upper limit is shown.

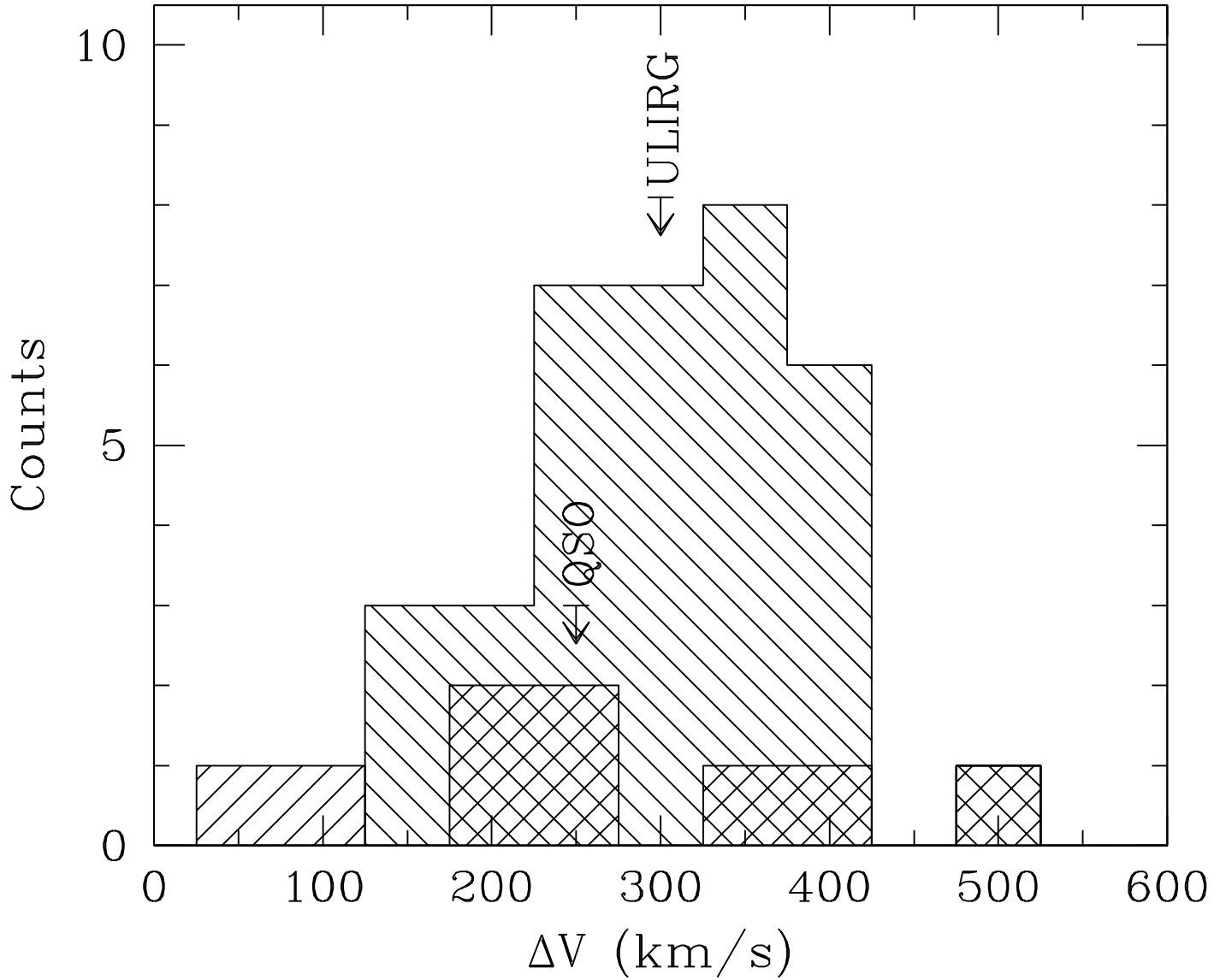


FIG. 7.— A histogram of CO line widths for the optically selected QSOs and ULIRGs. The sources plotted are the same as in Figure 8. The median line widths for the QSOs and the ULIRGs are also marked.

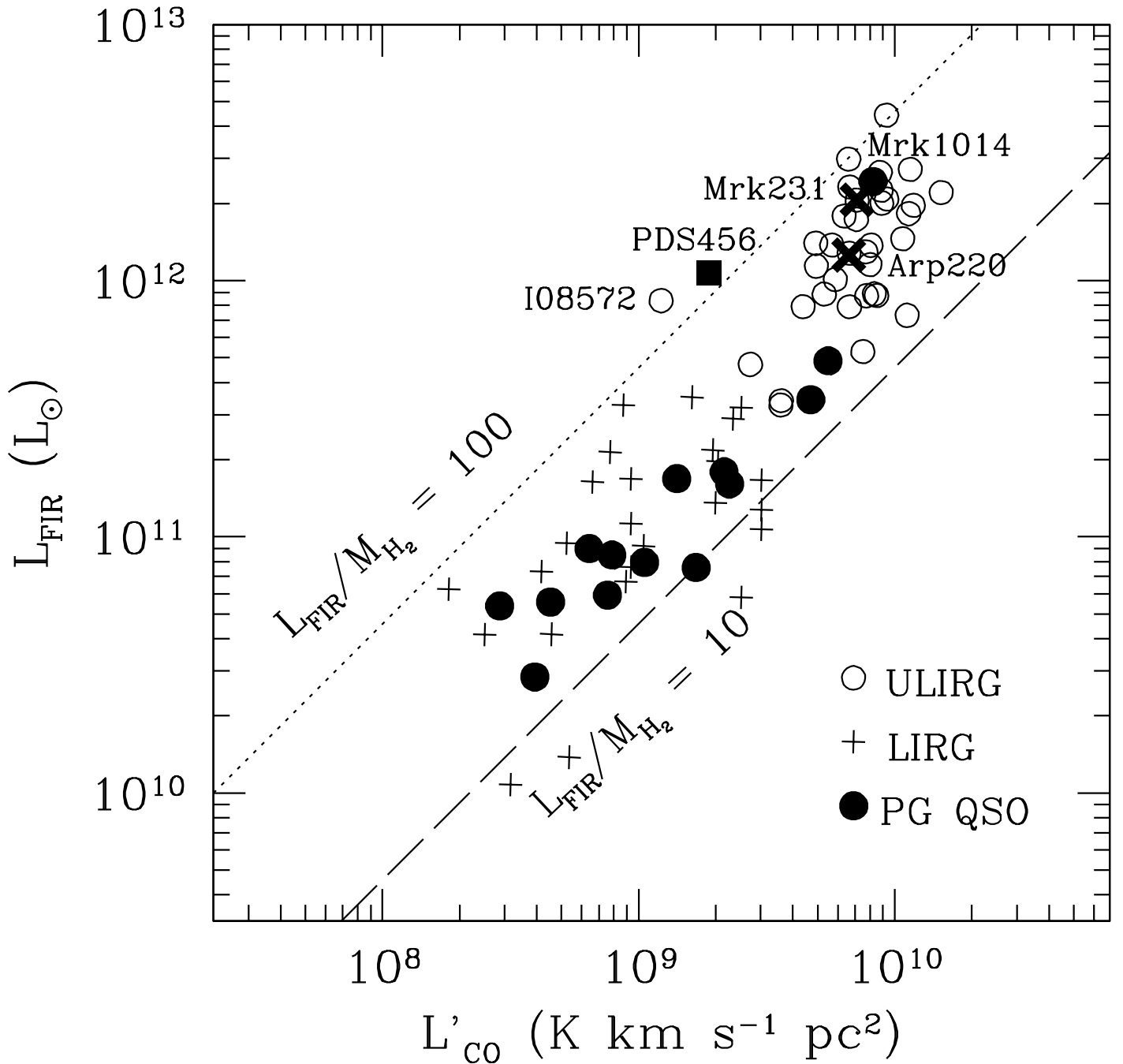


FIG. 8.— FIR luminosity versus L'_{CO} . PDS 456 is shown as a filled square while 14 optically selected PG QSOs with available CO measurements (Alloin et al. 1992; Evans et al. 2001; Scoville et al. 2003) are shown in filled circles. Typical uncertainty in all of these measurements are about 20%, and the size of the errorbars are comparable or smaller than the size of the symbols shown. The prototype ULIRG Arp 220 and IR QSO/ULIRG Mrk 231 are identified using a thick X. Luminous infrared galaxies (crosses; Sanders et al. 1991) and ULIRGs (empty circles; Solomon et al. 1997) are also plotted for comparison. L_{FIR} is derived using the definition given in Table 1 of Sanders & Mirabel (1996). L'_{CO} can be converted to H_2 mass (M_{H_2}) using the standard CO-to- H_2 conversion given by Young & Scoville (1991). High temperature and density in intense starburst and circum-AGN regions may cause an over-estimate of the gas mass by a factor of a few (see Scoville, Yun, & Bryant 1997; Downes & Solomon 1998), but this effect should apply uniformly to all objects compared here.

TABLE 1
SUMMARY OF THE *K*-BAND SOURCES IN THE PDS 456 FIELD

Object	α (J2000)	δ (J2000)	m_K
QSO	17:28:19.78	-14:15:56.12	9.89 ± 0.02
K1	17:28:19.64	-14:15:57.54	16.3 ± 0.2
K2+K3	17:28:19.69	-14:15:59.02	15.6 ± 0.1

TABLE 2
SPECTRAL ENERGY DISTRIBUTION DATA FOR PDS 456

λ	ν (GHz)	S_ν (mJy)	Reference
25 cm	1.2	24 ± 5	this work
21 cm	1.4	30.4	1
21 cm	1.4	22.7 ± 3.3	2
6 cm	4.85	8.23	1
3.6 cm	8.5	4.6 ± 0.9	this work
3000 μm	100	< 0.6	this work
850 μm	350	< 7.5	this work
100 μm	3000	420 ± 140	1
100 μm	3000	< 882	3
80 μm	3750	510 ± 110	1
60 μm	5000	930 ± 69	3
50 μm	6000	610 ± 110	1
25 μm	12000	750 ± 41	3
25 μm	12000	550 ± 120	1
12 μm	25000	514 ± 24	3
12 μm	25000	280 ± 30	1
4.8 μm	62500	225 ± 36	4
3.8 μm	78900	180 ± 23	4
2.2 μm	136400	78 ± 6	4
2.2 μm	136400	73 ± 7	this work
2.159 μm	139000	78.0 ± 0.8	5
1.662 μm	178300	38.6 ± 0.6	5
1.65 μm	181800	39 ± 3	4
1.25 μm	240000	26 ± 2	4
1.235 μm	242800	24.8 ± 0.3	5

References. — (1) Reeves et al. 2000; (2) Condon et al. 1998; (3) IRAS ADDSCAN; (4) Simpson et al. 1999; (5) 2MASS